The Quest for Tensor Modes

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Window of Opportunity

Measurement of sizeable tensor modes could give precious information on the physics at highest energies.

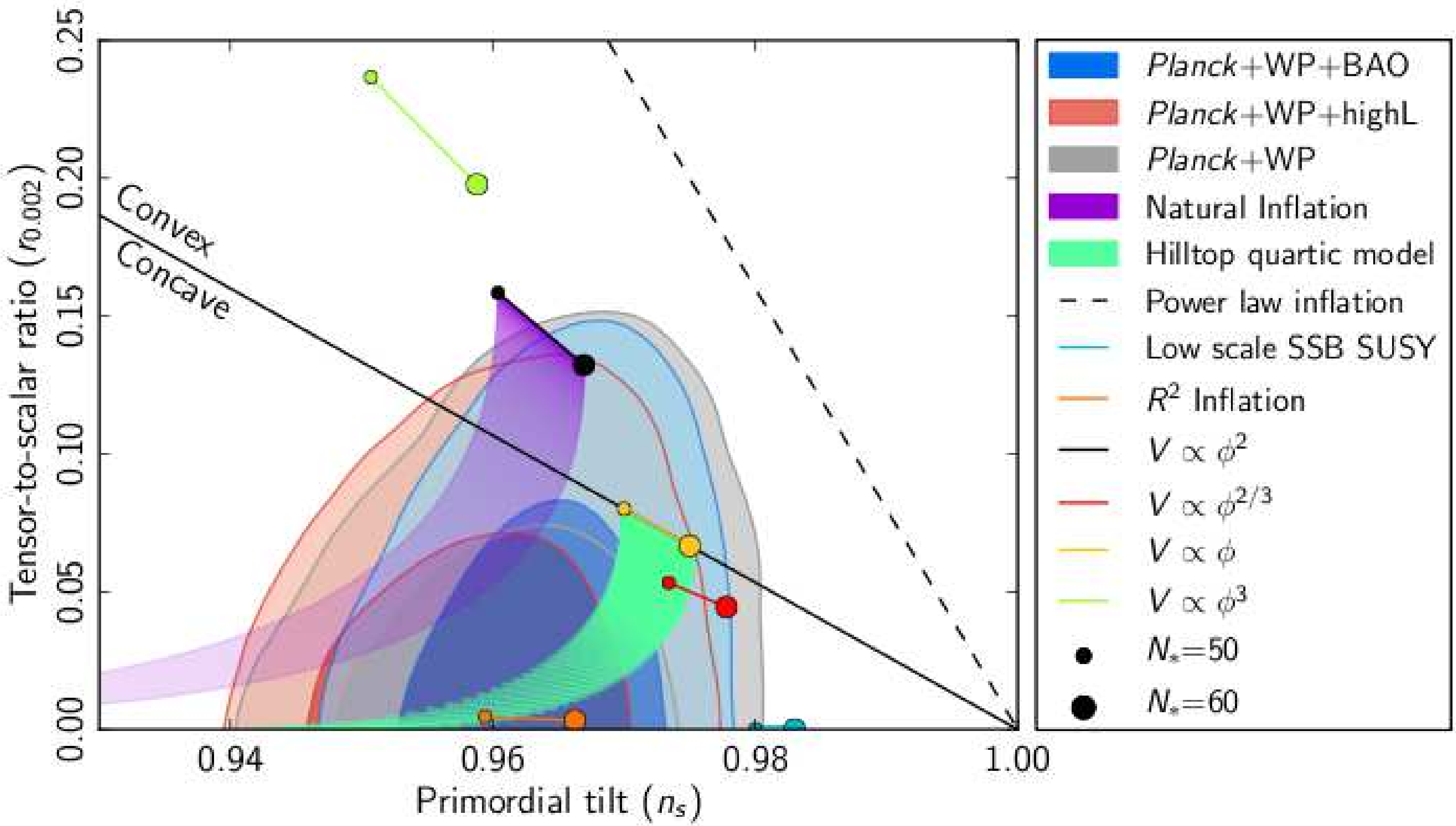
- we might explore physics close to the Planck scale
- get new insight on properties of (quantum) gravity

A theoretical treatment of these sizeable tensor modes could therefore be problematic

- question of trans-Planckian excursions (Lyth bound)
- the question of UV-completion (e.g. string theory)
- the weak gravity conjecture (and variants thereof)

Let us hope that sizeable tensor modes will be found.

Planck results (Spring 2013)



Outline

We consider here two specific cases for large tensor modes

- the mechanism of axionic (natural) inflation

  (Freese, Frieman, Olinto, 1990)

- enhanced tensor modes from extra dimensions

  (Giudice, Kolb, Lesgourgues, Riotto, 2002)

On the way we shall be confronted with some obstacles

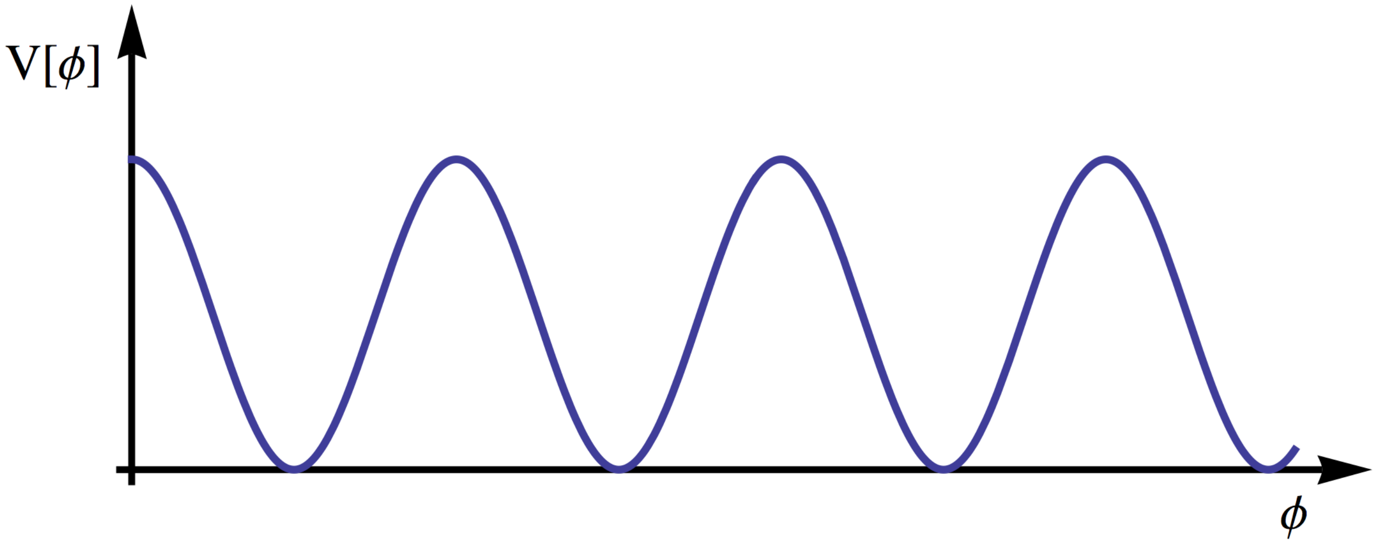
- the "problem" of trans-Planckian decay constants
- duality symmetries in string theory
- upper limits on Hubble scale during inflation

Based on work by (Kappl, Nilles, Winkler, 2015; Im, Nilles, Trautner, 2017)

Axionic Inflation

The axion exhibits a shift symmetry $\phi \to \phi + c$

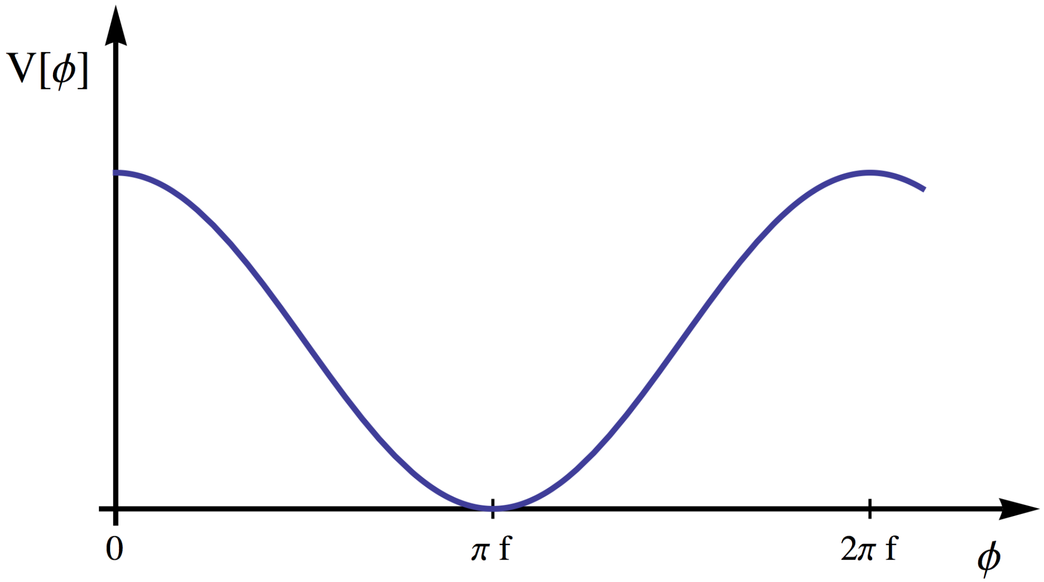
Nonperturbative effects break this symmetry to a remnant discrete shift symmetry



$$V(\phi) = \Lambda^4 \left[1 + \cos\left(\frac{\phi}{f}\right)\right]$$

The Axion Potential

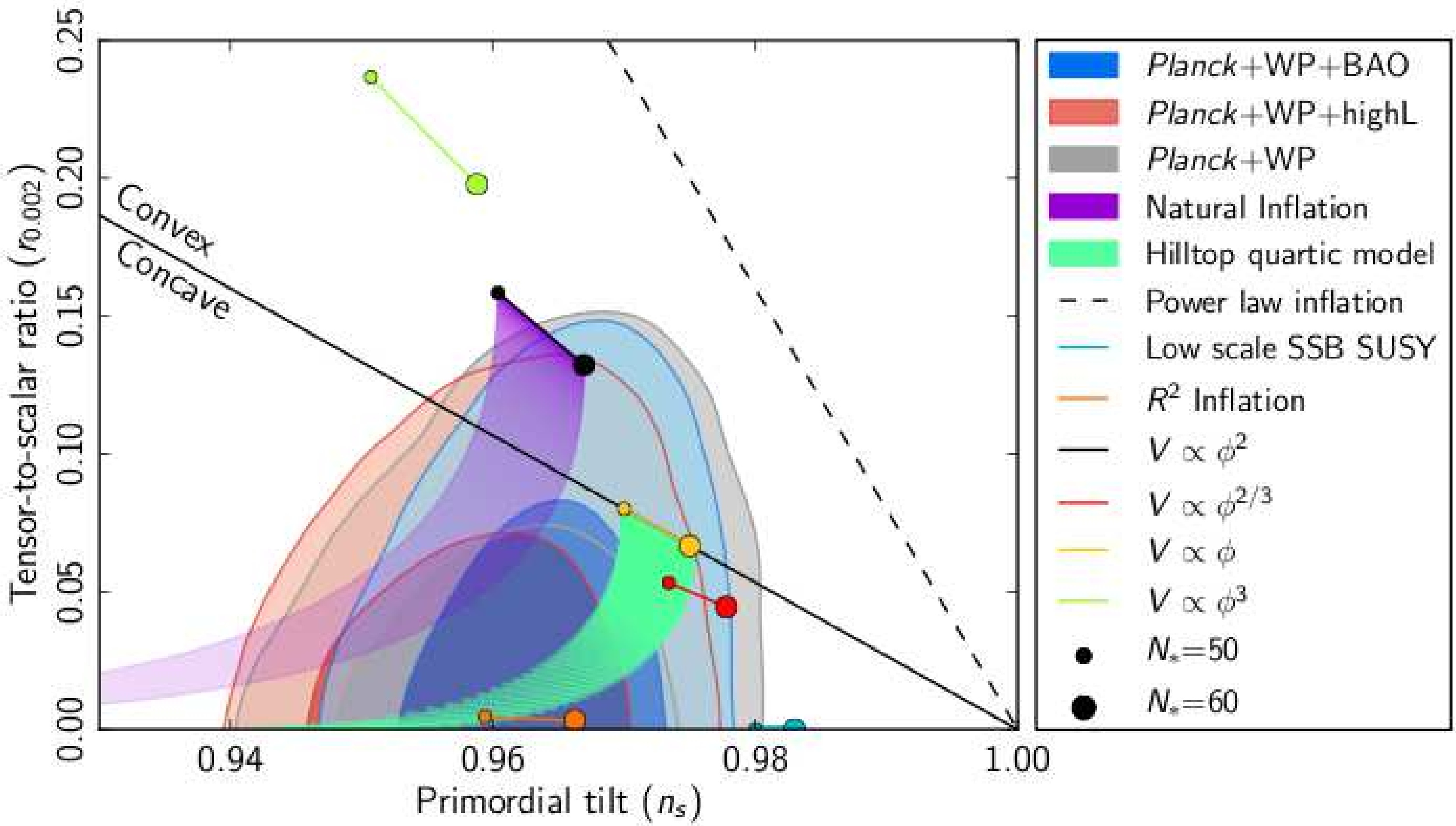
Discrete shift symmetry identifies $\phi = \phi + 2\pi n f$



$$V(\phi) = \Lambda^4 \left[1 + \cos\left(\frac{\phi}{f}\right) \right]$$

ϕ confined to one fundamental domain

Planck results (Spring 2013)



The question of tensor modes

Sizeable tensor modes are of particular interest:

- they might lead us to scales of physics close to the Planck scale and the so-called “Lyth bound”
- potential $V(\phi)$ of order of GUT scale
- trans-Planckian excursions of the inflaton field
- For a quadratic potential $V(\phi) \sim m^2\phi^2$ it implies $\triangle\phi \sim 15 M_{\mathrm{P}}$ to obtain 60 e-folds of inflation

Axionic inflation, on the other hand, seems to require the decay constant to be limited: $f \leq M_{\mathrm{P}}$

So this might be problematic, in particular for a UV-completion in string theory (Banks, Dine, Fox, Gorbatov, 2003)

Aligned axions

A way out is the consideration of two (or more) fields.

(Kim, Nilles, Peloso, 2004)

- still we require $f \leq M_{\rm P}$ for the individual axions
- aligned axion has effective decay constant beyond the Planck scale
- string favours a multi-axion picture

The alignment prolongs the fundamental domain of the aligned axion to super-Planckian values, as the axionic inflaton spirals down in the potential of the second axion

This avoids the restrictions of the single axion model.

The KNP set-up

We consider two axions

$$\mathcal{L}(\theta,\rho) = (\partial\theta)^2 + (\partial\rho)^2 - V(\rho,\theta)$$

with potential

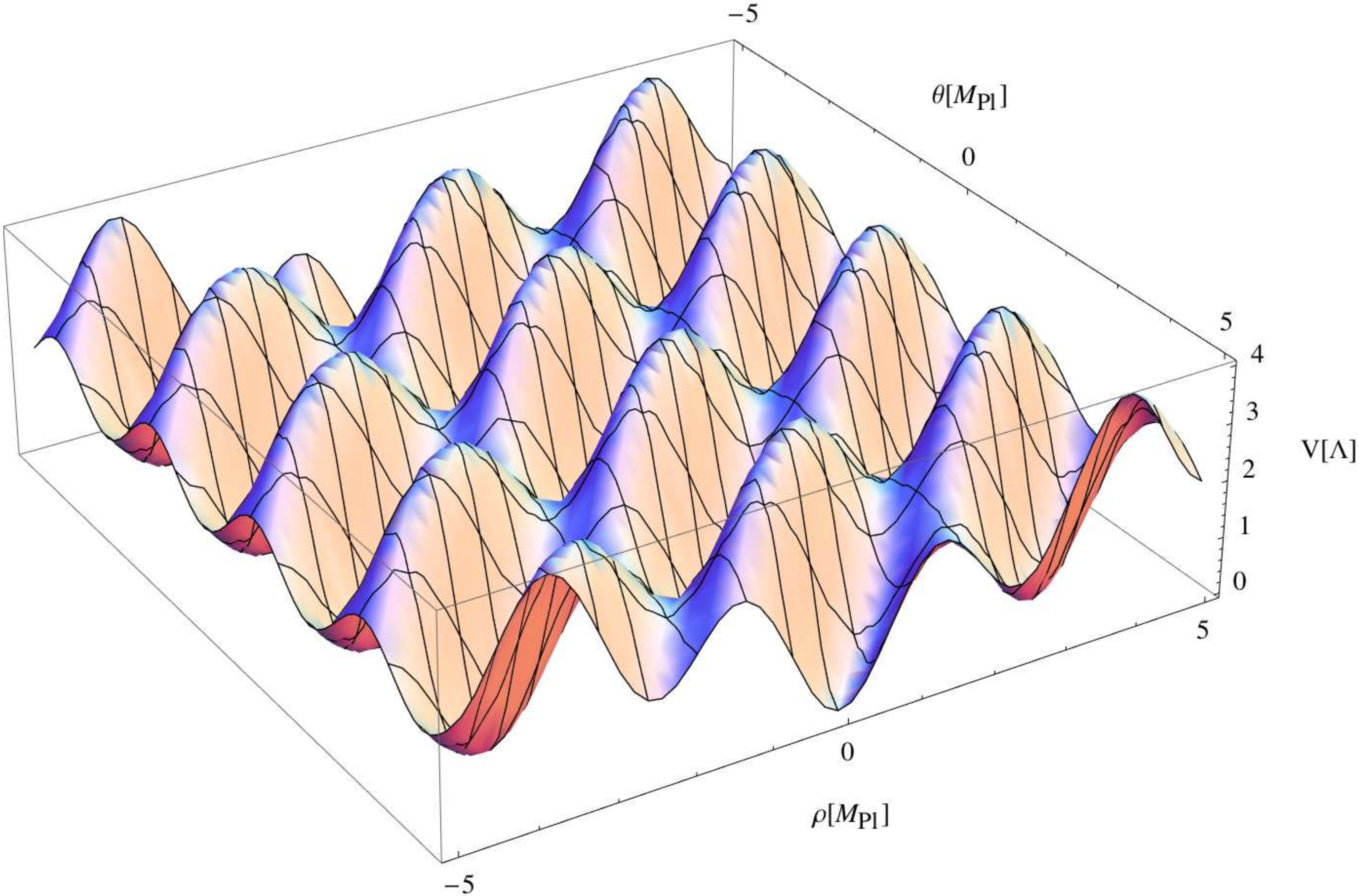
$$V(\theta,\rho) = \Lambda^4\left(2 - \cos\left(\frac{\theta}{f_1} + \frac{\rho}{g_1}\right) - \cos\left(\frac{\theta}{f_2} + \frac{\rho}{g_2}\right)\right)$$

This potential has a flat direction if $\frac{f_1}{g_1} = \frac{f_2}{g_2}$

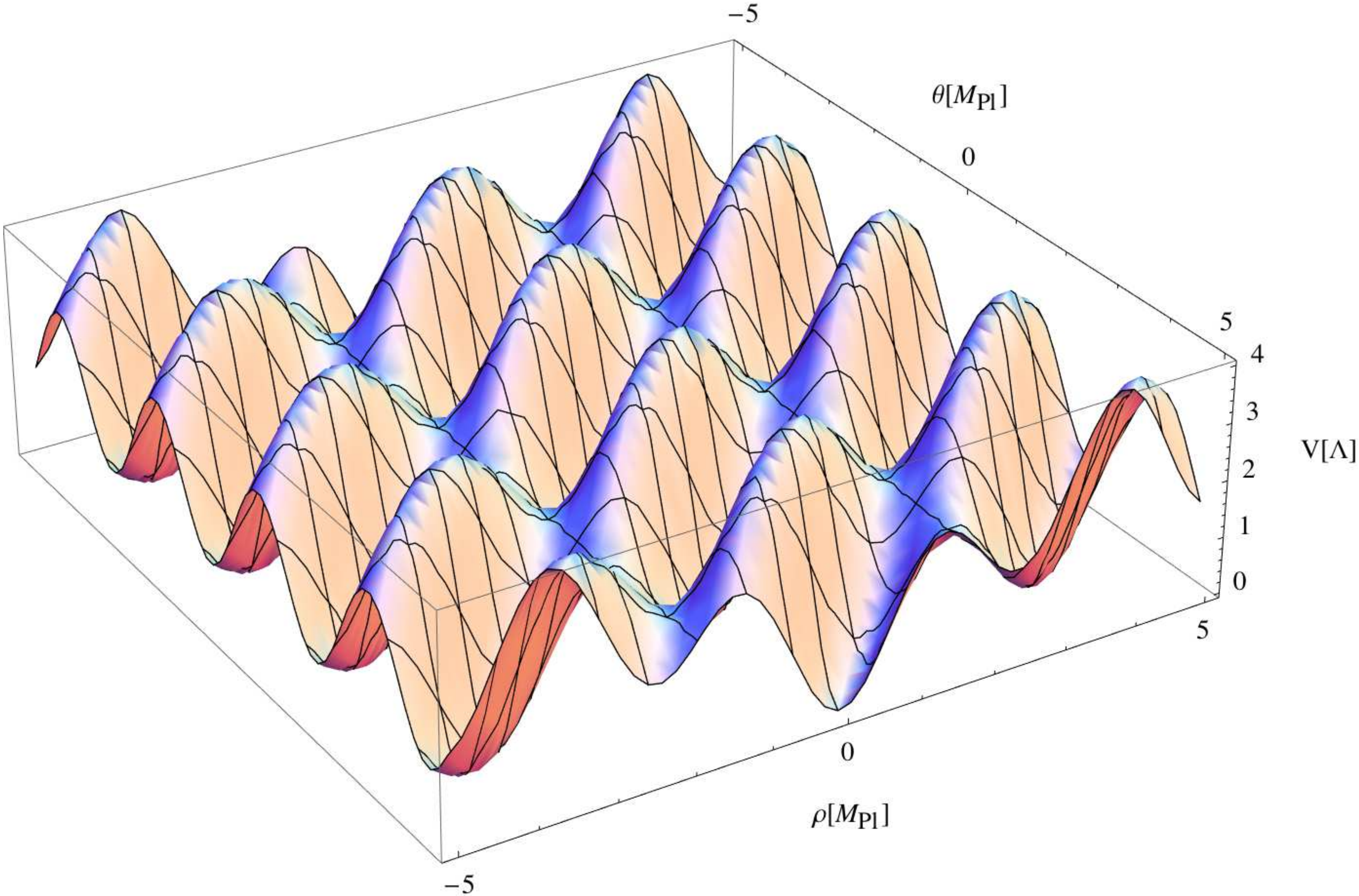
Alignment parameter defined through $\alpha = g_2 - \frac{f_2}{f_1} g_1$

For $\alpha = 0$ we have a massless field ξ.

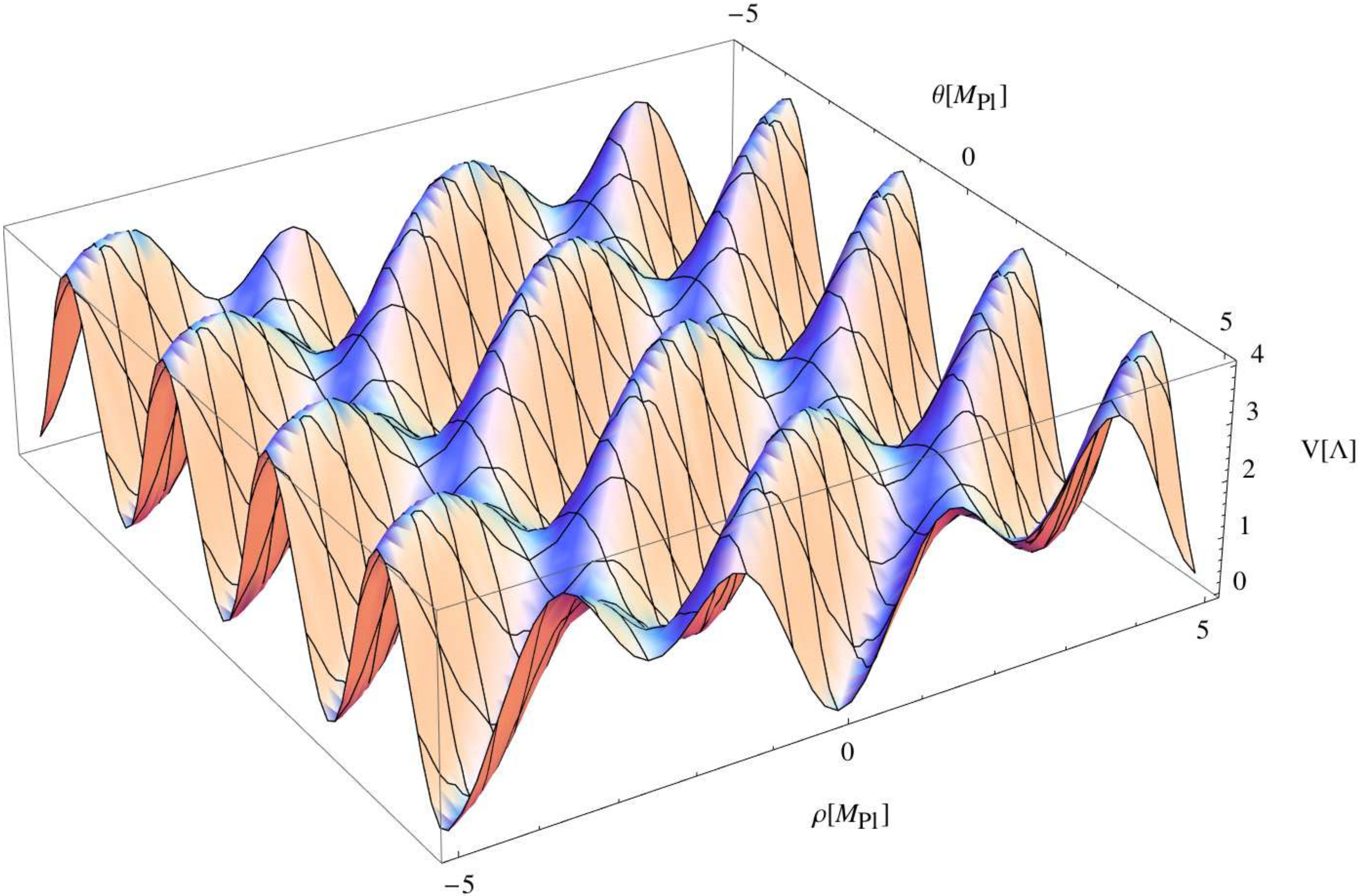
Potential for $\alpha = 1.0$



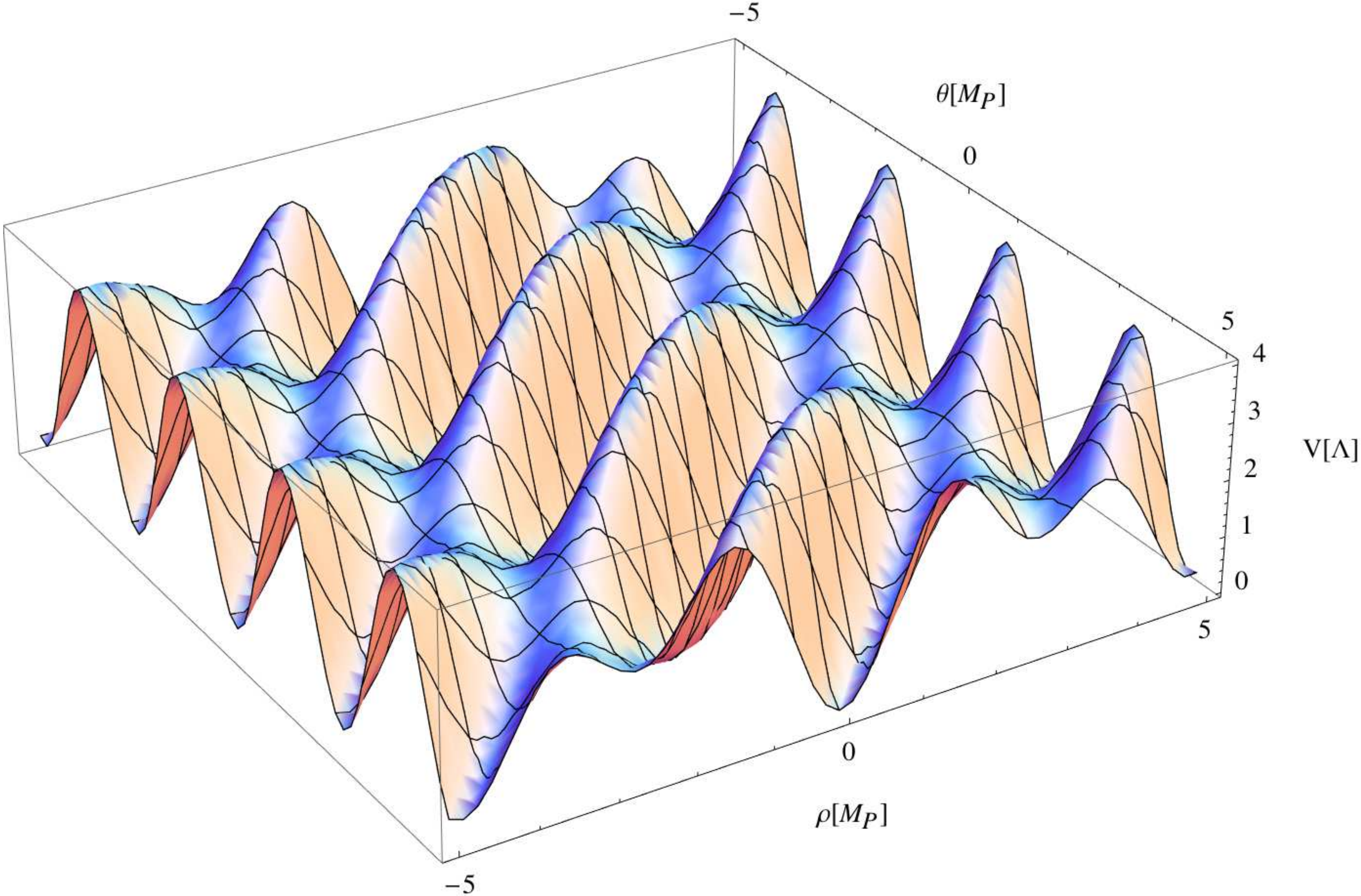
Potential for $\alpha = 0.8$



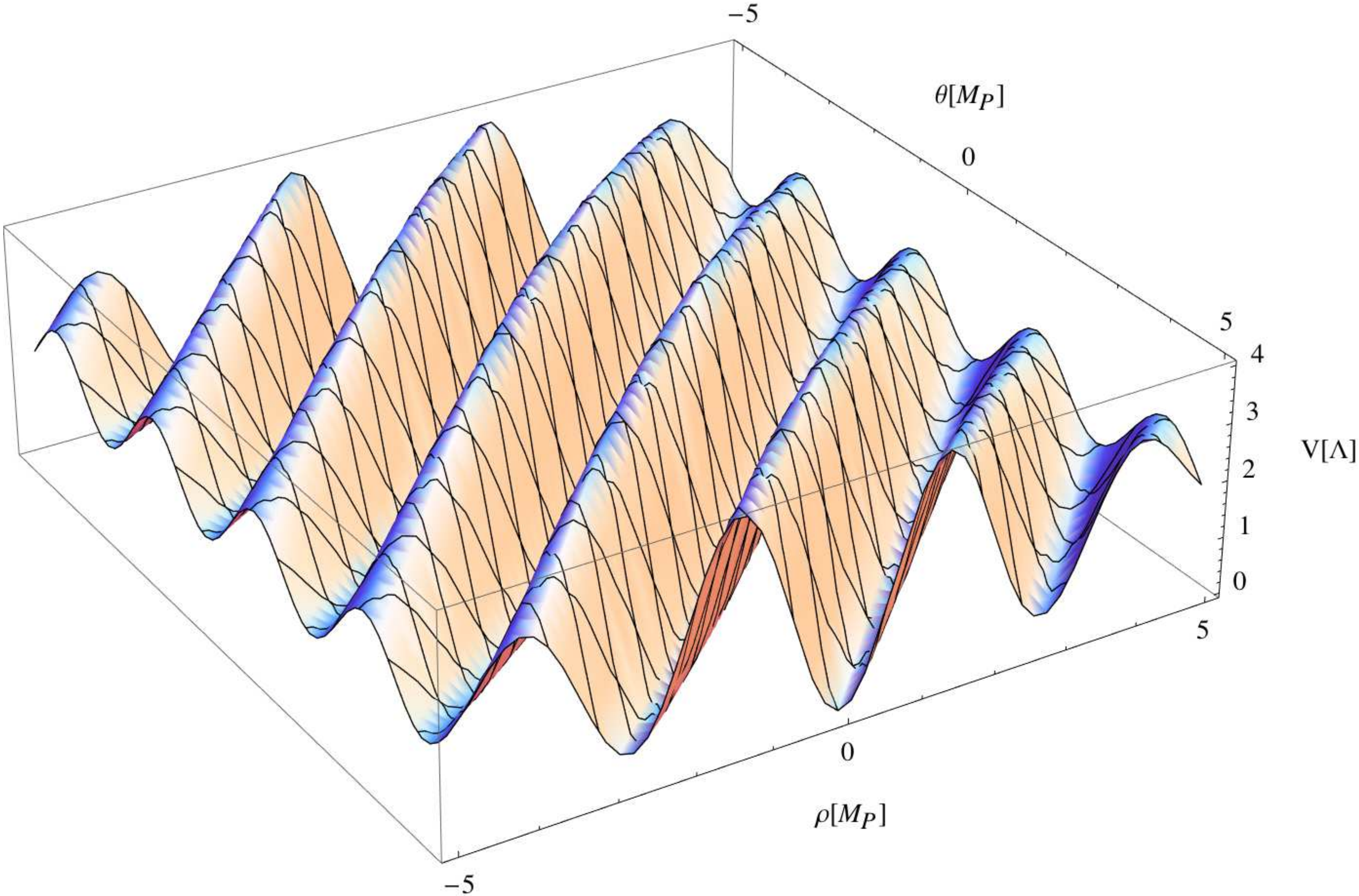
Potential for $\alpha = 0.5$



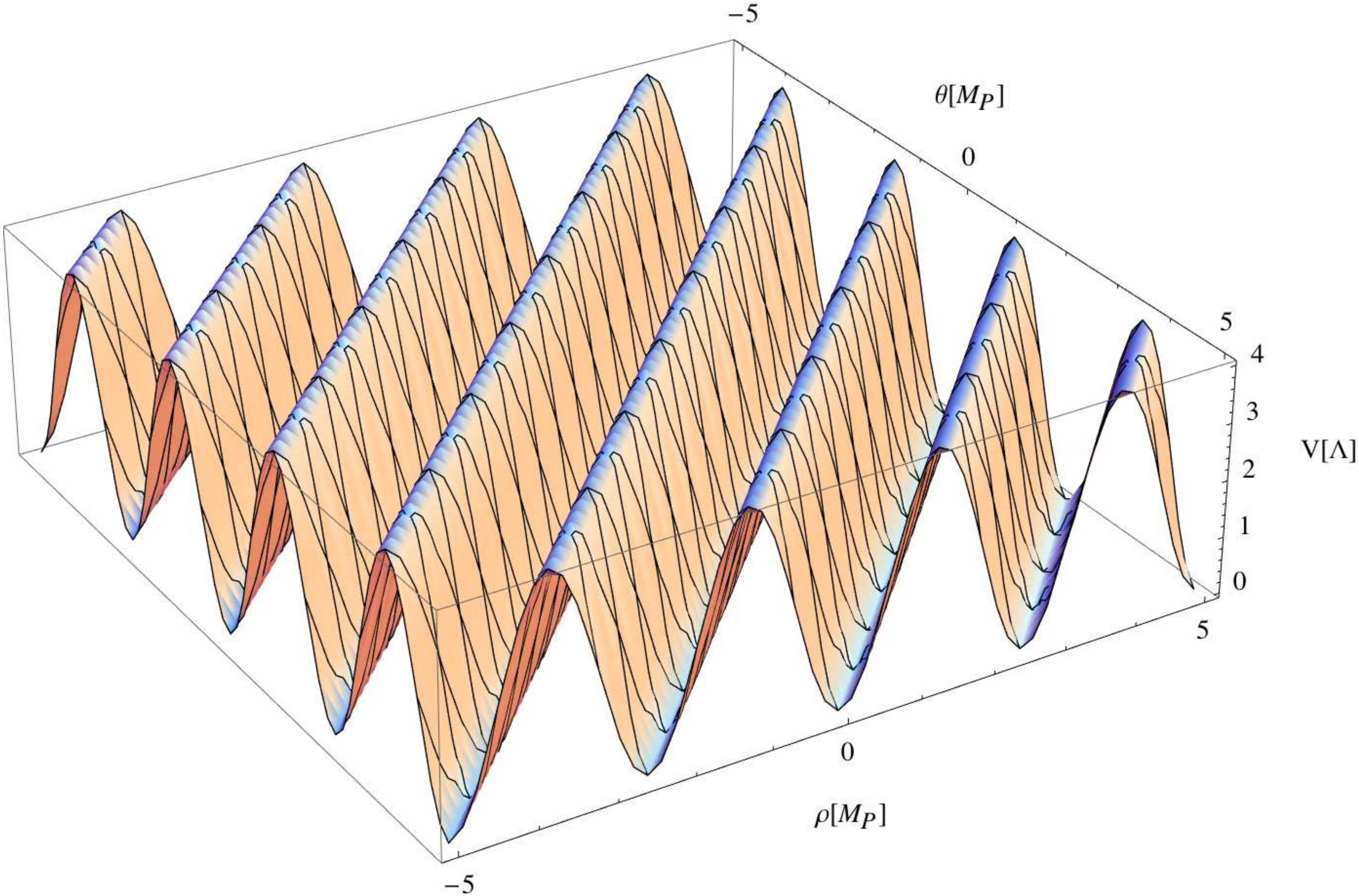
Potential for $\alpha = 0.3$



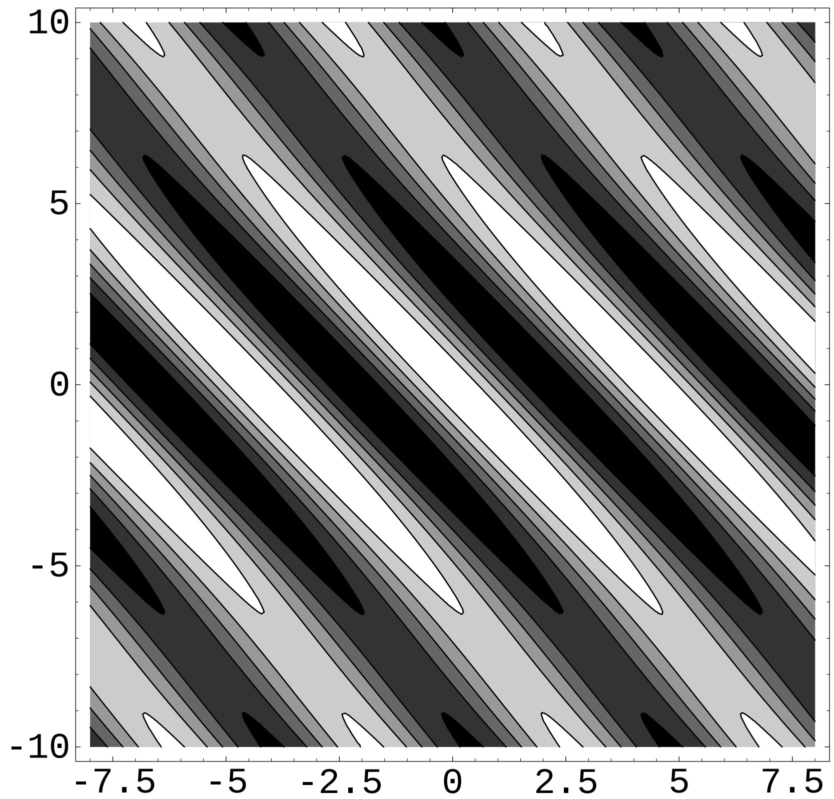
Potential for $\alpha = 0.1$



Potential for $\alpha = 0$



The aligned Axion Landscape



The field ξ rolls within the valley of ψ. The motion of ξ corresponds to a motion of θ and ρ over many cycles. The system is still controlled by discrete symmetries.

UV-Stability

We have a very flat direction and within the effective QFT we are at the “edge of control”

- is inflation perturbed by other effects?
- is there an upper limit on $f_{\rm eff}$?

Remember that in case of a single axion we had the limit

- $f \leq M_{\rm string}$ (Banks, Dine, Fox, Gorbatov, 2003)
- derived from dualities in string theory (e.g. T-duality)

In the multi-axion case these arguments are not directly applicable, but we still have to worry about these questions, also in view of the "weak gravity conjecture".

(Arkani-Hamed, Motl, Nicolis, Vafa, 2006)

$T-$Duality

String dualities give important constraints on the axion decay constants, especially $T-$duality $SL(2,Z)$:

$$T \to \frac{aT - ib}{icT + d}$$

generated by an inversion and a shift

$$T \to 1/T, \qquad T \to T + i.$$

$$G = K + \log|W|^2$$

must be invariant under T-duality.

$T-$**Duality**

K and W might transform nontrivially. Consider e.g.

$$K = -3 \log (T + \bar{T}).$$

This Kähler potential transforms under $SL(2, Z)$ as

$$K \to K + \log |icT + d|^6$$

and has to be compensated by a superpotential transforming as a modular form (of weight -3):

$$W \to (icT + d)^{-3} W \ .$$

Not just a Cosine

In string theory we do not just get cosine potentials, but obtain modular functions (e.g. Dedekind-functions) from

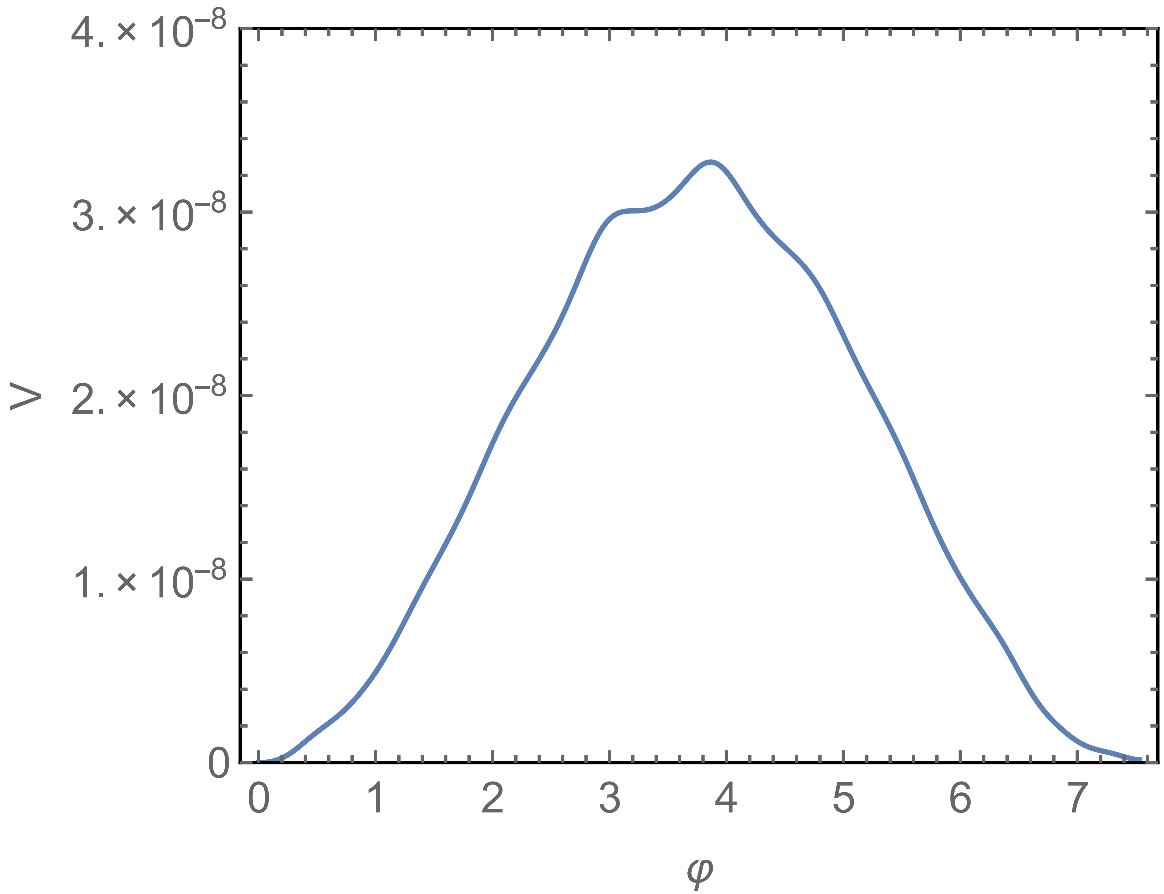
- world sheet instanton effects,
- gauge kinetic functions and gaugino condensates.

So we might consider instead a modular function

$$\eta(T) = e^{-\pi T/12} \times \prod_k \left(1 - e^{-2k\pi T}\right)$$

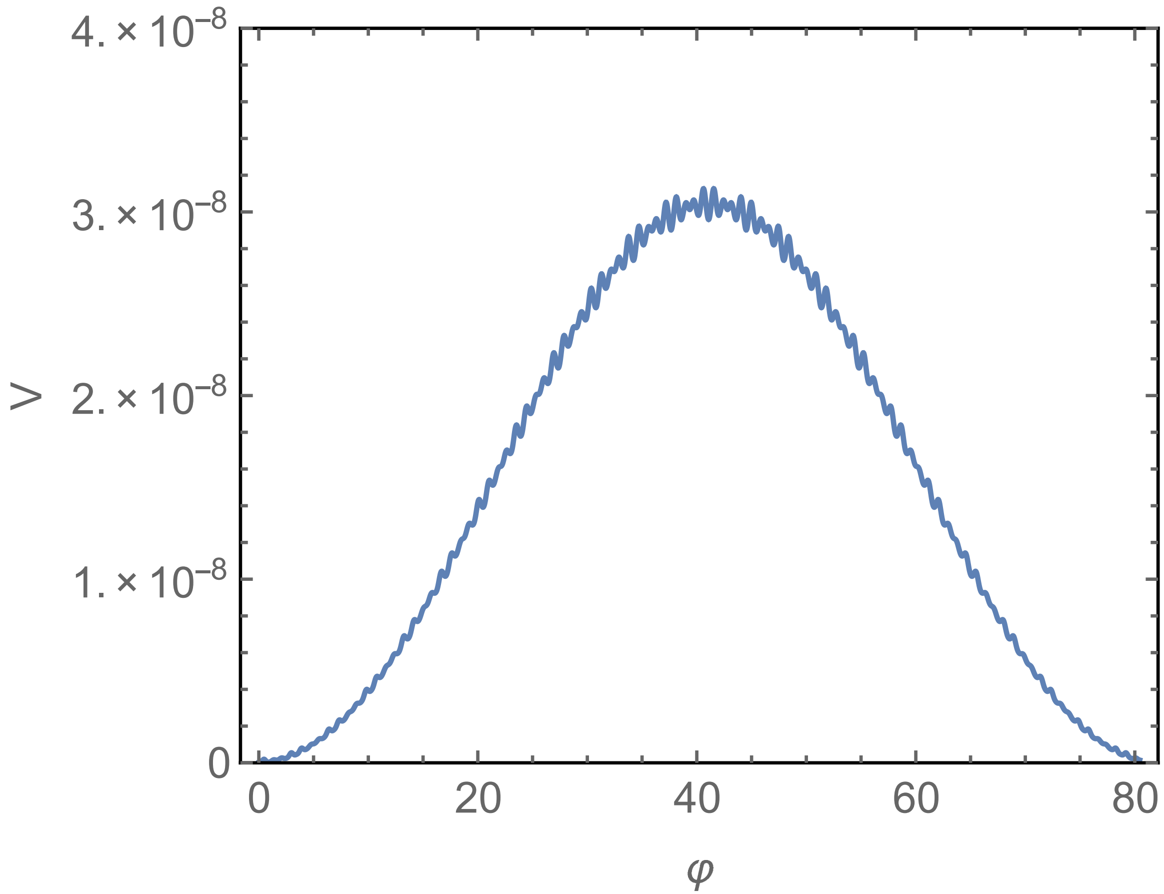
The higher harmonics give wiggles in the potential that perturb the flat direction and might stop inflation. In the case of a single axion this prevents a trans-Planckian f.

Wiggles in the aligned potential



The wiggles in the case of weak alignment (small f)

Wiggles in the aligned potential



Strong alignment (leading to superPlanckian f)

Modulated natural Inflation

It seems plausible that under some circumstances the wiggles become important and

- spoil the flat direction,
- provide un upper limit on decay constant $f_{\rm eff}$.

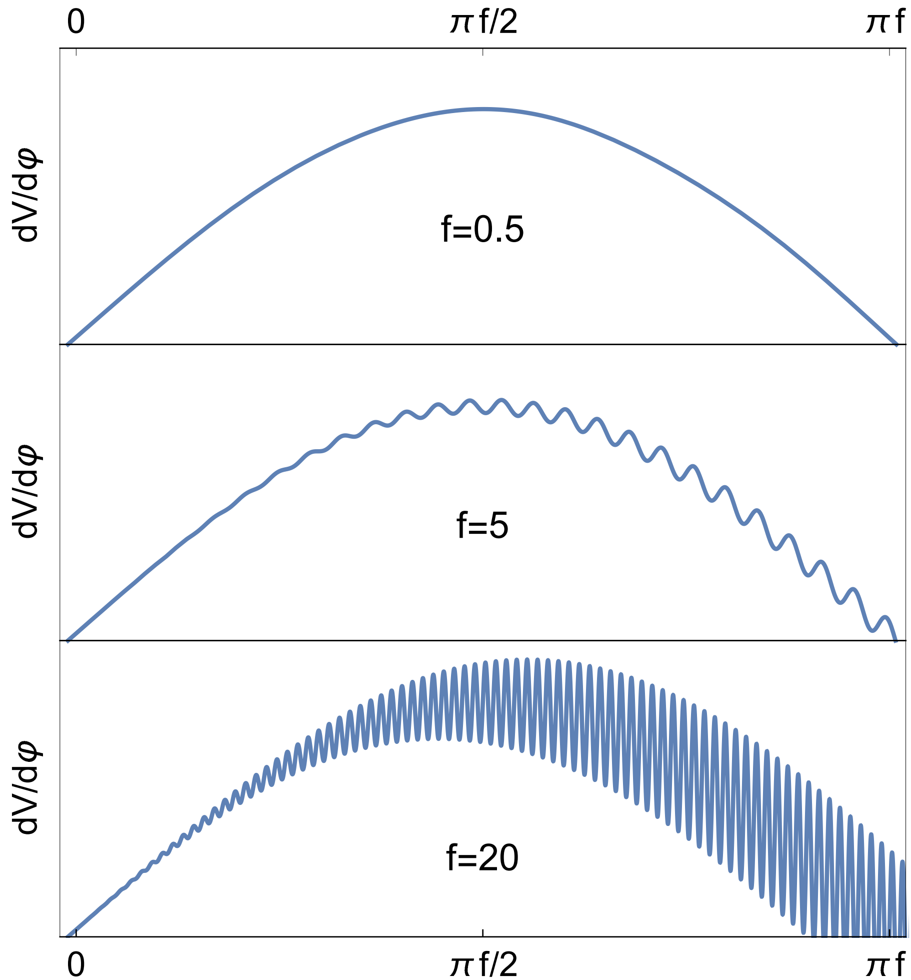
Explicit calculation are necessary to clarify the situation,

- but might be beyond our present capabilities;
- observational confirmation is extremely important.

Restrictions from WGC are satisfied here both in the aligned and non-aligned case.

(Kappl, Nilles, Winkler, 2015; Choi, Kim, 2015; Kobayashi, Nitta, Urakawa, 2016)

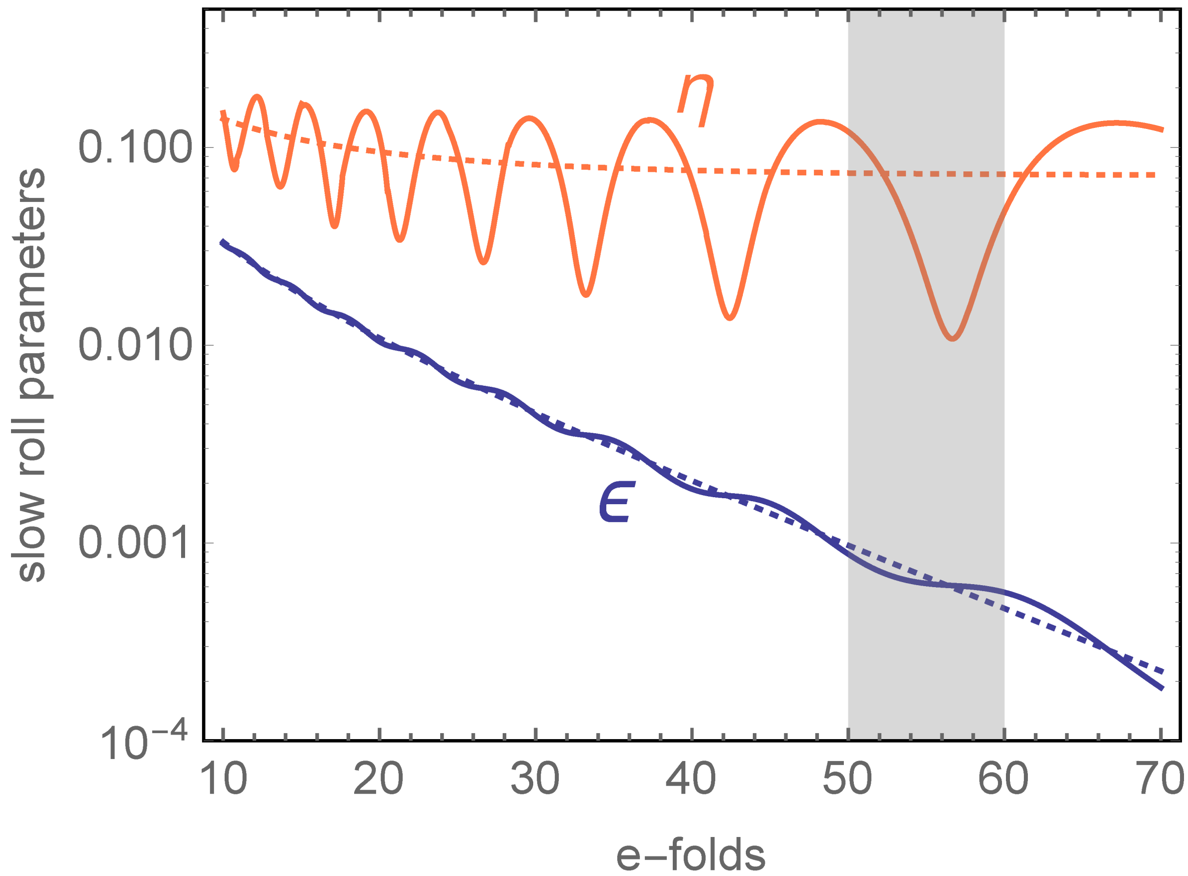
Slope of Potential



Wiggly structure becomes important

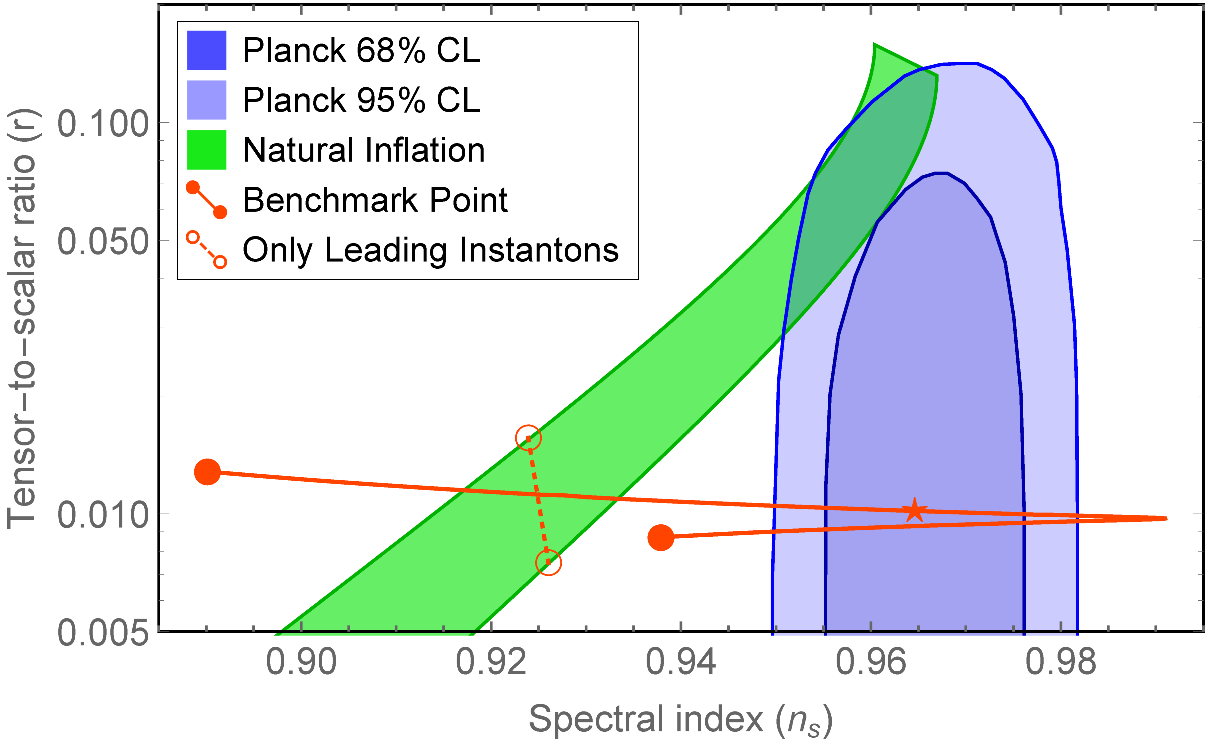
(Abe, Kobayashi, Otsuka, 2015; Kappl, Nilles, Winkler, 2015)

Slow roll parameters



ϵ (η) depend on first (second) derivative

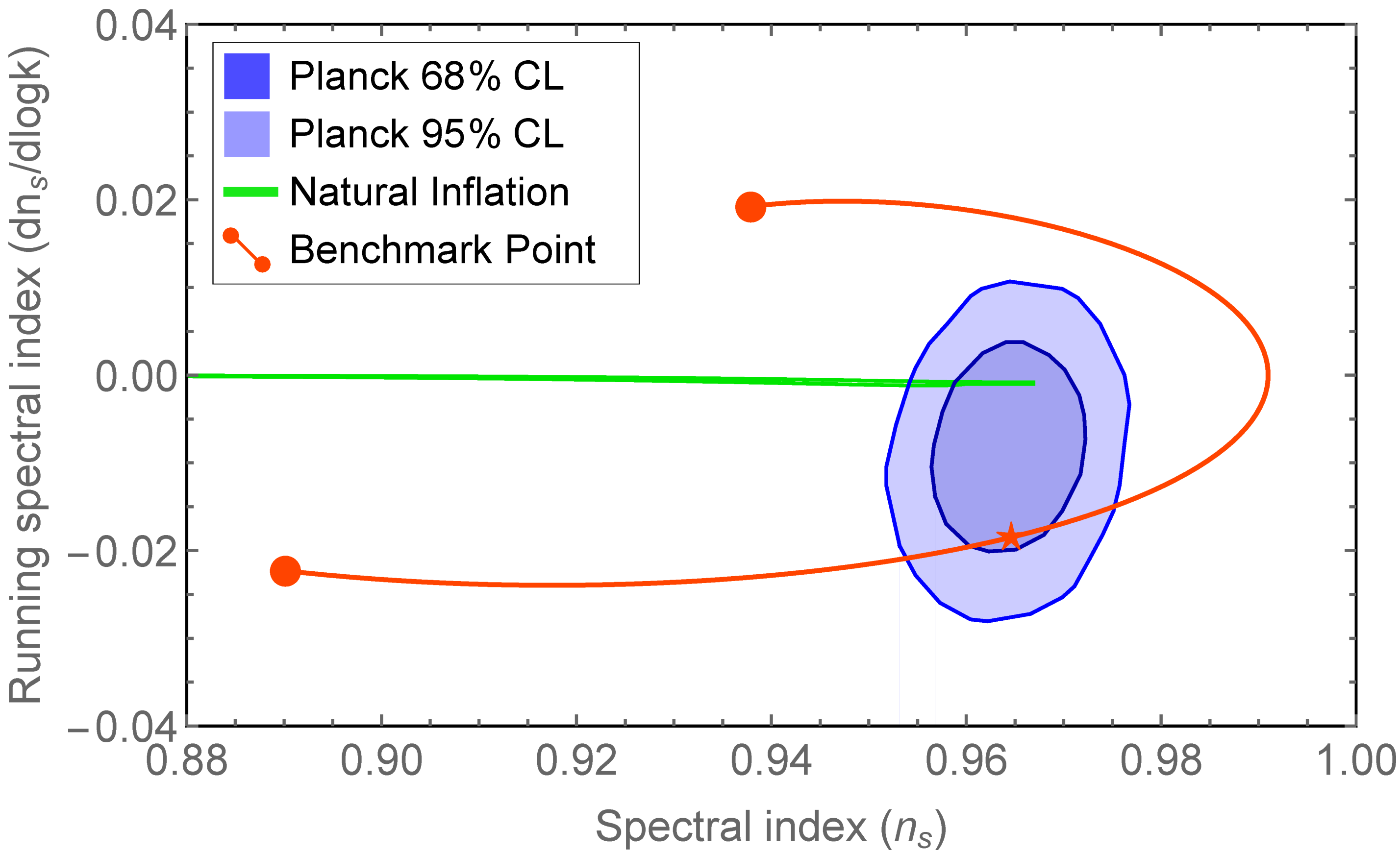
$n_s - r$ plane



Strong variations of n_s on the number of e-folds.

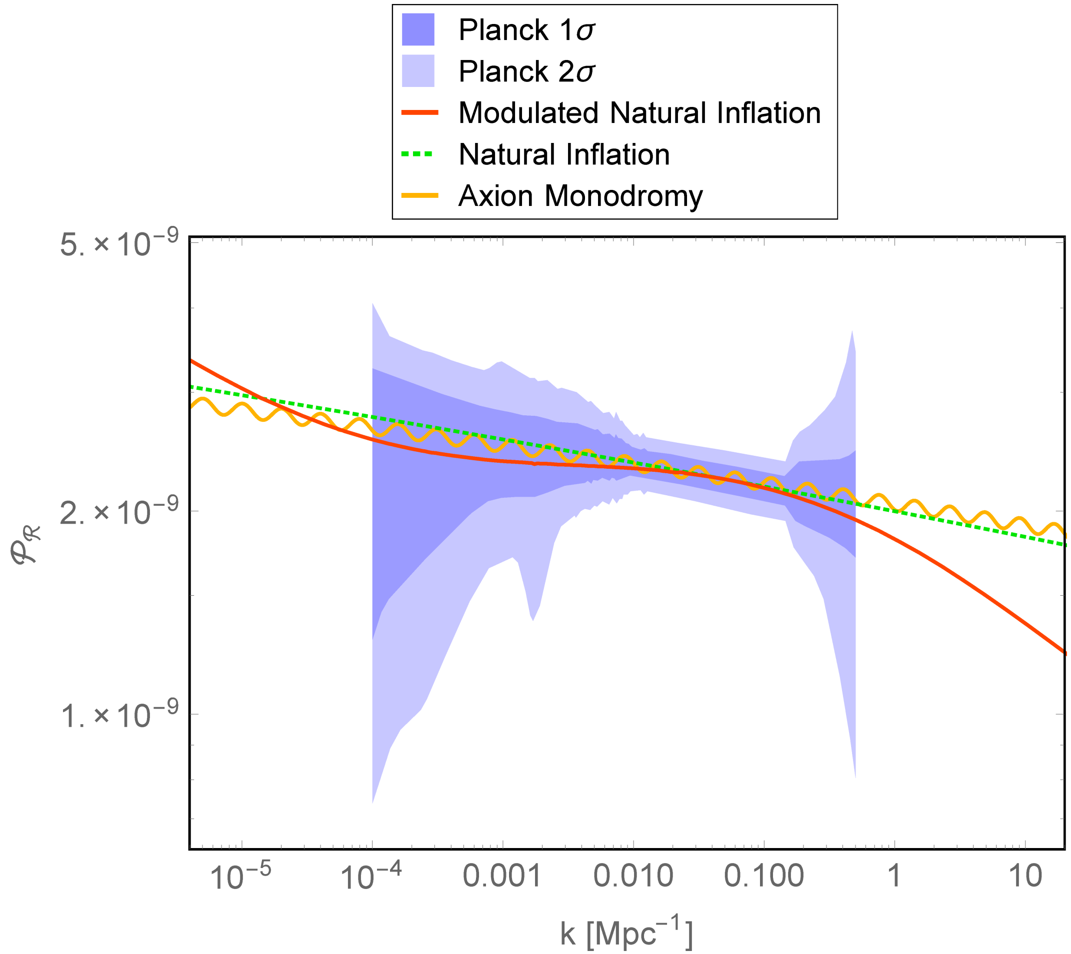
(Kappl, Nilles, Winkler, 2015)

Running of spectral index



Comparison of spectral index with Planck data

Scalar power spectrum



Comparison to Planck reconstructed power spectrum

Intermediate Summary

Sizeable tensor modes are possible,

- but there is most probably an upper limit on $f_{\rm eff}$ and r,
- we might have "wiggles" and running indices

We are in a region where we lose computational control, at least in the traditional 4-dimensional set-up.

- the UV-completion requires higher harmonics
- the "cosine" is no longer just a "cosine"
- we need experimental observations to clarify the situation and teach us new lessons about gravity.

Still we have to explore possible alternatives.

Extra dimensions

Extra space dimensions might provide new ingredients:

- strength of gravity could vary in bulk
- weak scale hierarchy problem might be solved via large or warped extra dimensions

This could influence the size of tensor modes. We consider

- large extra dimensions (LED) (Arkani-H., Dimopoulos, Dvali, 1998)
- warped extra dimensions (RS) (Randall, Sundrum, 1999)
- linear dilaton model (LD) (Antoniadis, Dimopoulos, Giveon, 2001)

Matter fields live on visible brane, gravity on hidden brane. The inflaton can reside on either of them. (Im, Nilles, Trautner, 2017)

The mechanism

We start with a warm-up example

- one large extra dimension (LED)
- inflaton field on visible (IR) brane
- we assume that the mechanism of radius stabilisation does not influence the tensor modes

This simplification is called IRB assumption

- scalar modes are essentially d=4 dimensional

(Giudice, Kolb, Lesgourgues, Riotto, 2002)

- tensor modes are influenced by bulk effects
- relevance of effective Planck mass during inflation

Effective Planck Mass

Effective Planck mass during inflation

$$M^2_{\mathrm{Pl,eff}} \;=\; M^3\, 2\,\pi\, R\, \left(1 - \tfrac{2}{3}\,\pi^2\, R^2\, H^2\right)$$

has to be compared to

$$M^2_{\mathrm{Pl,eff}}\Big|_{H=0}^{\mathrm{LED}} = \; M^3\, 2\,\pi\, R\;.$$

This leads to enhanced tensor modes and

$$\tfrac{2}{3}\,\pi^2\, R^2\, H^2 \;<\; 1$$

implies an upper limit on Hubble scale during inflation.

(Im, Nilles, Trautner, 2017)

Results

Results are model dependent (Im, Nilles, Trautner, 2017)

- we obtain enhanced tensor modes in LED and RS within IRB-scenario
- surprisingly the IRB assumption is not applicable to the Linear Dilaton (LD) case
- a satisfactory picture requires contributions from the hidden brane ("remote inflation")
- "remote inflation" is an option in the RS scenario as well

  (Nihei, 1999; Kaloper, 1999; Kim, Kim, 1999)
- LD case with dilaton as stabilizer field is consistent with IRB assumption and leads to reduced tensor modes

  (Kehagias, Riotto, 2016)

Summary

Sizeable tensor modes are a challenge for model building

- especially for a consistent UV-completion
- restrictions from string theory via dualities
- effective theories (like aligned axions) as one way out
  - expect upper limit on $f_{\rm eff}$ (how large can it be?)
  - specific signatures of "modulated natural inflation"
- extra dimensions as another way out
  - enhanced tensor modes imply upper limit on H
  - remote inflation as a new option
  - implication of reduced tensor modes?

Sizeable tensor modes: a window of opportunity