Collider signals and neutralino dark matter detection in relic-density-consistent models without soft term universality

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Neutralino Dark Matter

- Dominant composition of matter in our universe is not detected visibly but inferred from gravitational effects (Galactic Clustering, Rotation Curves, Gravitational Lensing, Cosmic Microwave Background ...)

- Dark Matter should be non-baryonic (no candidate in the SM), non-relativistic (cold), stable(or long-lived), weakly (or super-weakly) interacting matter

- In SUSY models with R-parity conservation
  $\Rightarrow$ the Lightest Supersymmetric Particle(LSP) is absolutely stable
  $\Rightarrow$ lightest neutralino $\tilde{Z}_1$ is the LSP in most of MSSM parameter space
  $\Longrightarrow$ $\tilde{Z}_1$ is good candidate for Cold Dark Matter (CDM)

- From the WMAP results, the cold dark matter density of the universe is $\Omega_{CDM}h^2 = 0.111^{+0.011}_{-0.015}$: (upper bound is a tight constraint on SUSY models containing DM candidates : DM may consist of several components)

- Number density is governed by Boltzmann equation,
  $dn/dt = -3Hn - \langle \sigma v_{rel} \rangle (n^2 - n_0^2)$
  $\Rightarrow$ requires evaluating many thousands Feynman diagrams
  $\Longrightarrow$ high (co-)annihilation cross section implies low relic abundance

Motivations for SUSY models with non-universal soft terms

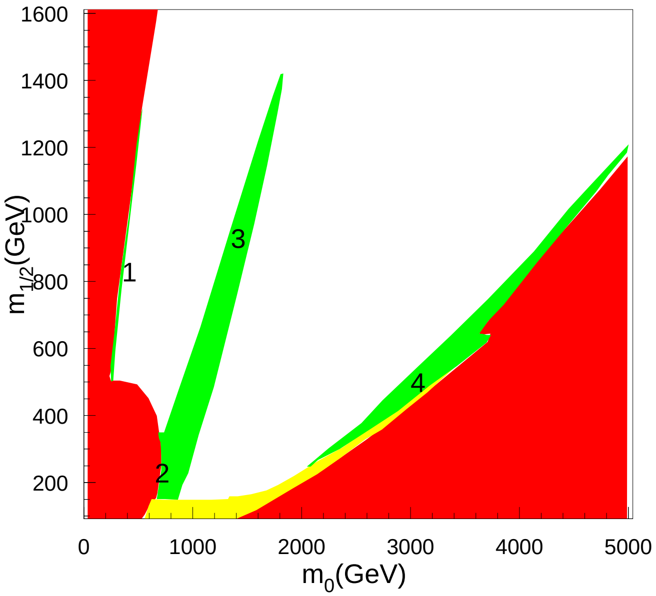
- **WMAP allowed Regions** in parameter space of mSUSGRA

  1. $\tilde{\tau}$ co-annihilation region at low m_0, $m_{\tilde{\tau}_1} \sim m_{\tilde{Z}_1}$

  2. bulk region at low m_0 and $m_{1/2}$, light sleptons (LEP2 excluded)

  3. Higgs-funnel H, A resonance($2m_{\tilde{Z}_1} \simeq m_{A,H}$) at large tan$\beta \sim 50$ or h-resonance at low $m_{1/2}$ ($2m_{\tilde{Z}_1} \simeq m_h$)

  4. FP/HB region at large m_0, low $\mu \rightarrow$ mixed higgsino dark matter (MHDM)



- **Motivations for models with non-universality**

  - ⋆ all relic-density-consistent regions in mSUGRA are near the edges of theoretically(or experimentally) excluded
  - ⋆ need to examine how already drawn conclusions from the mSUGRA model are affected by relaxing the universality assumptions
  - ⋆ within R-parity conserved neutralino dark matter assumption, WMAP value provides a strong constraint reducing model parameter space by one unit

Assumptions and Tools

- **Assumptions**
  - MSSM is an effective theory between the weak and GUT scale
  - R-parity is conserved
  - Neutralino LSP
  - (near)degeneracy of first and second generation of SSB sfermions → FCNC suppressed
  - CP-violating phases in SSB suppressed → CP contribution of SUSY is small
- **Tools**
  - all mass spectrum: ISAJET 7.76
  - relic density, direct detection rate: IsaTools package (IsaReD, IsaReS)
  - all indirect dark matter detection rates: DarkSUSY
- **Relic-density-consistent models** obtained by adjusting
  - composition of neutralino (WTN: Well-Tempered Neutralino*)

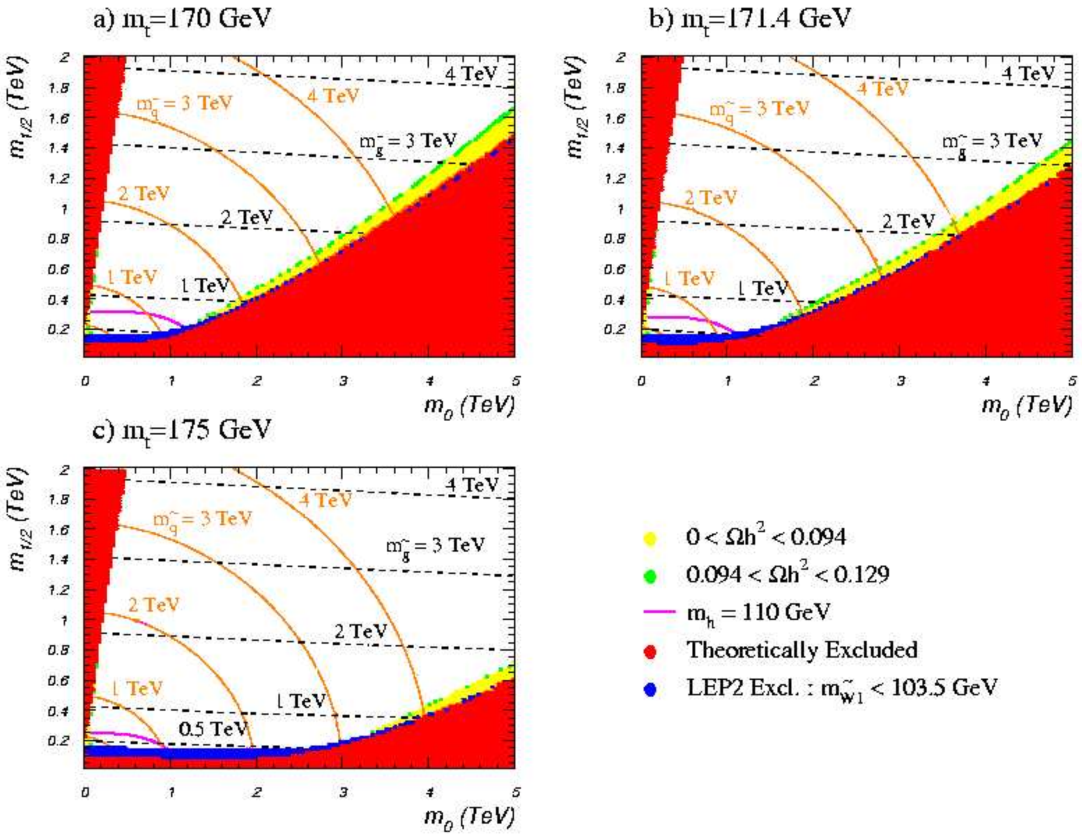
    *termed by Arkani-Hamed et al. Nucl.Phys.B741, 108, 2006
  - masses of neutralino or other sparticles

Updated results on parameter space of mSUGRA 1

- **parameter Space**

  m_0, $m_{1/2}$, A_0, $\tan\beta$, $sign(\mu)$
- **effect of m_t variation**
  - location of EWSB excluded region

    → HB/FP region moves
  - gluino and squark mass contours hardly change
  - as m_t increased, A-funnel moves to smaller m_0



mSUGRA : tanβ=10, A_0=0, μ>0

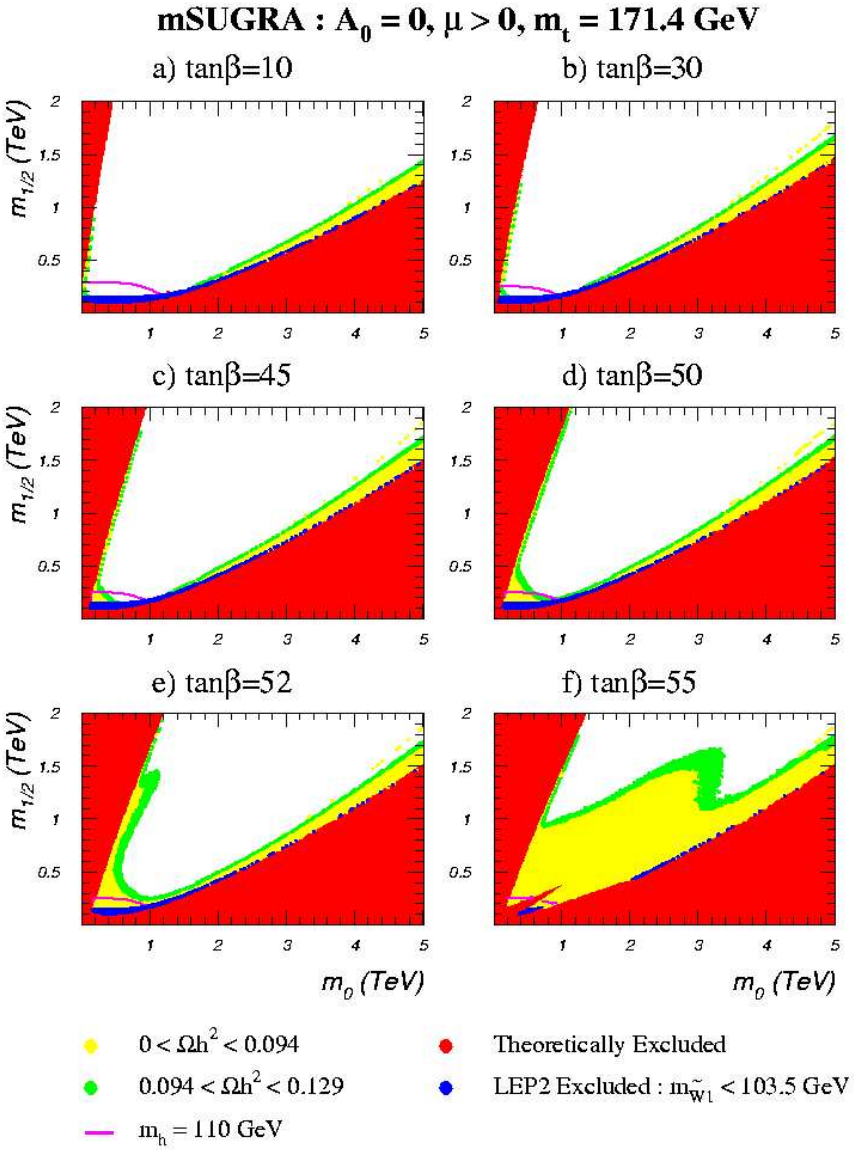
a) m_t=170 GeV

b) m_t=171.4 GeV

c) m_t=175 GeV

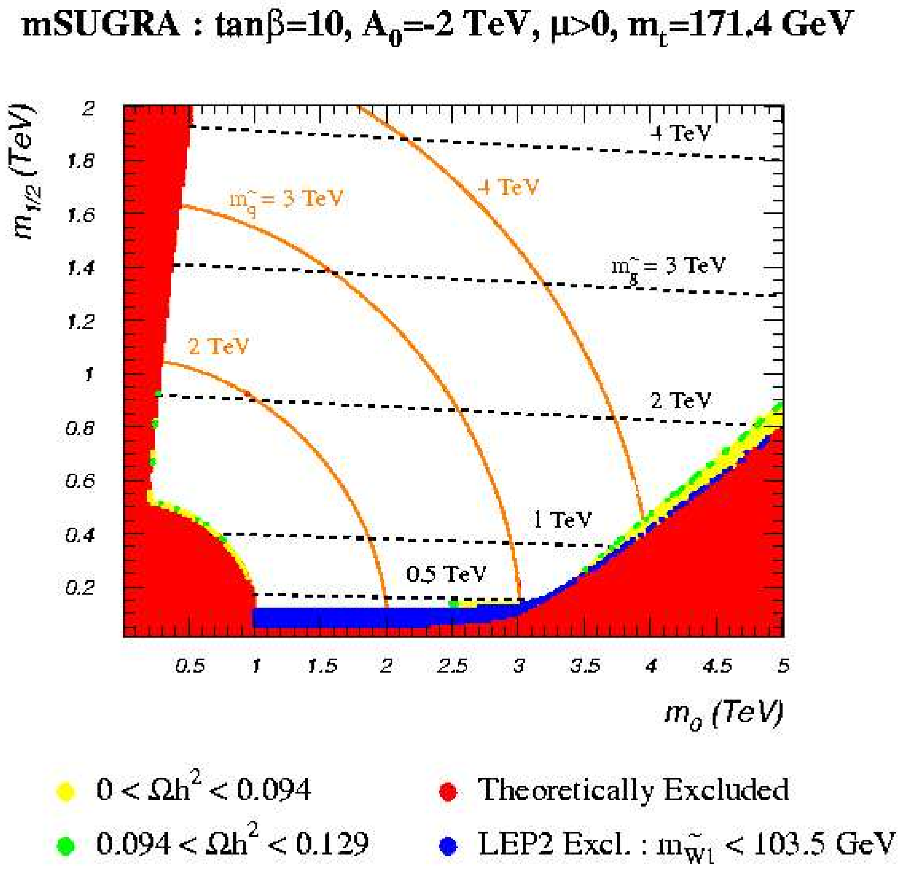
Updated results on parameter space of mSUGRA 2

- **effect of tanβ variation**
  - larger tanβ, wider WMAP allowed region
    (as tanβ grows, m_A drops due to b and τ Yukawa coupling effect)
  - as tanβ increased, A-funnel moves to larger m_0
- **effect of $\mu < 0$**
  - A-funnel arises at lower tanβ
  - narrower A-funnel



Updated results on parameter space of mSUGRA 3

- **effect of large negative A_0**
  $\rightarrow$ reduce $m_{\tilde{t}_1}$
  $\rightarrow$ $\tilde{t}_1 - \tilde{Z}_1$ mass gap quite small
  $\rightarrow$ relic density reduced via stop co-annihilation



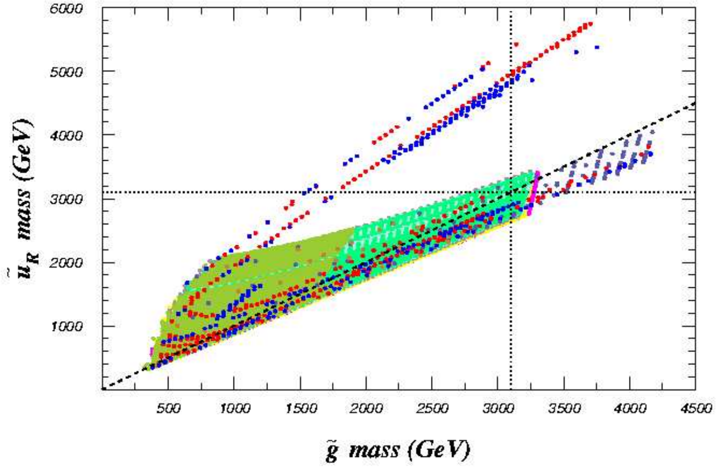
Non-universal scalar mass models

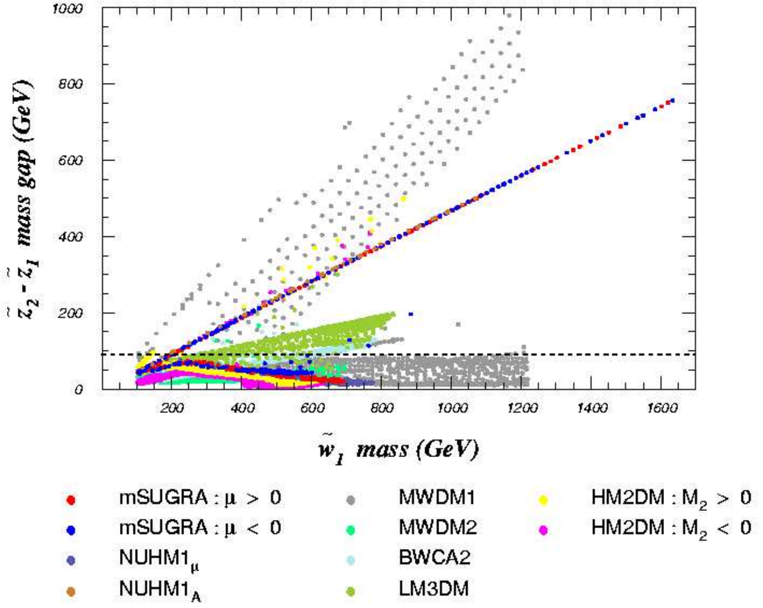
- generation non-universality: Normal scalar Mass Hierarchy (**NMH**)
  $m_0(1,2)$, m_0, $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - $m_0(1,2)$: first/second generation, $m_0(3) = m_{H_u} = m_{H_d} \equiv m_0$: remaining
  - dial $m_0(1,2)$ to low enough to bulk (co-)annihilation via light sleptons
- non-universal Higgs mass: one extra parameter case (**NUHM1**$_\mu$, **NUHM1**$_A$)
  m_0, δ_ϕ, $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - $m_\phi = m_0(1+\delta_\phi)$, $m^2_{H_u} = m^2_{H_d} \equiv sign(m_\phi)|m_\phi|^2$
  - $m_\phi > m_0$: small μ and MHDM
  - $m_\phi < 0$: $m_A \sim 2m_{\tilde{Z}_1}$ → at any tanβ
- non-universal Higgs mass: two extra parameter case (**HS**-Higgs Splitting)
  m_0, $m^2_{H_u}$(equivalently μ), $m^2_{H_d}$(equivalently m_A), $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - $m^2_{H_{u,d}} = m_0^2\ (1 \mp \delta_H)$
  - $\delta_H < 0$: low μ and low m_A
  - $\delta_H > 0$: WMAP region via $\tilde{l}_L/\tilde{\nu}$ or $\tilde{u}_R/\tilde{c}_R$ co-annihilation

Non-universal gaugino mass models

- Mixed Wino Dark Matter (**MWDM1**, **MWDM2**):
  m_0, M_1(or M_2), $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - by increasing the wino content of the LSP by reducing the ratio M_2/M_1
  - $M_1 \neq M_2 = M_3 = m_{1/2}$ or $M_2 \neq M_1 = M_3 = m_{1/2}$
- Bino-Wino Co-Annihilation Scenario (BWCA1, BWCA2):
  same as MWDM but M_1 and M_2 are in opposite sign
  - by allowing co-annihilation between high bino-like and wino-like states
- Low $|M_3|$ Dark Matter: Compressed SUSY (**LM3DM**):
  m_0, M_3, $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - by increasing the higgsino content of the LSP by decreasing the gluino mass
  - $M_3 \neq M_1 = M_2 = m_{1/2}$
- High $|M_2|$ Dark Matter: left-right split SUSY (**HM2DM**):
  m_0, M_2, $m_{1/2}$, A_0, tanβ, $sign(\mu)$
  - by allowing large M_2 mass
  - $M_2 >> M_1 = M_3 = m_{1/2}$

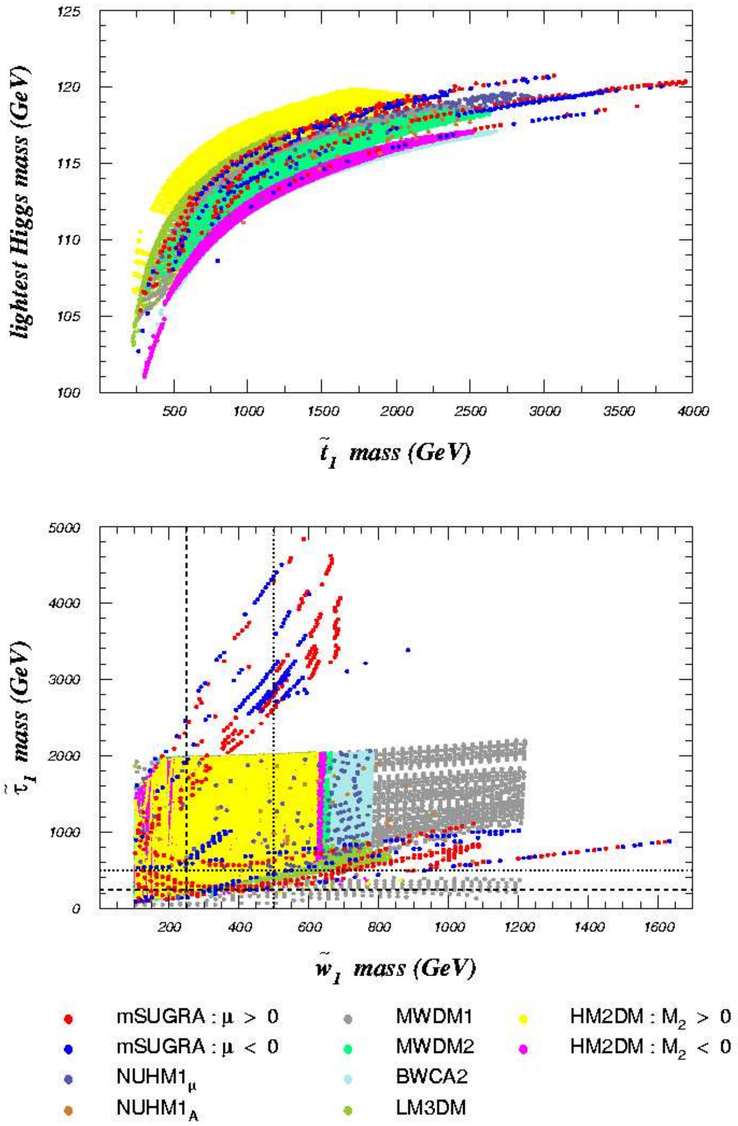
Implications for collider searches 1





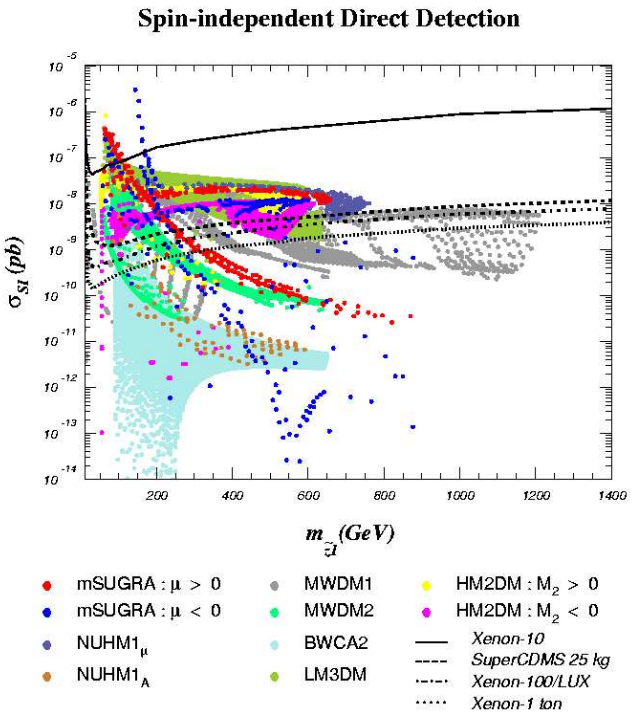
- – with $A_0 = 0$, $m_t = 171.4$ GeV, $\tan\beta = 10$ (except for the mSUGRA model: $\tan\beta = 10$, 30, 45, 50, 52 and 55)
  - – non-universal mass dialed to yield $\Omega_{\widetilde{Z}_1} h^2 \simeq 0.11$
- $m_{\tilde{g}}$ *vs.* $m_{\tilde{u}_R}$
  - – dotted lines: 100 fb^{-1} reach of CERN LHC
  - – dashed line: $m_{\tilde{u}_R} = m_{\tilde{g}}$
  - – most of models within reach of LHC except HB/FP region of mSUGRA
- $m_{\widetilde{W}_1}$ *vs.* $m_{\widetilde{Z}_2} - m_{\widetilde{Z}_1}$
  - – dashed line: $m_{\widetilde{Z}_2} - m_{\widetilde{Z}_1} = M_Z$
  - – below the line, 3-body decay like $\tilde{Z}_2 \to \tilde{Z}_1 l\bar{l}$ open
  - – in most models, $m(l\bar{l})$ mass edge visible at LHC

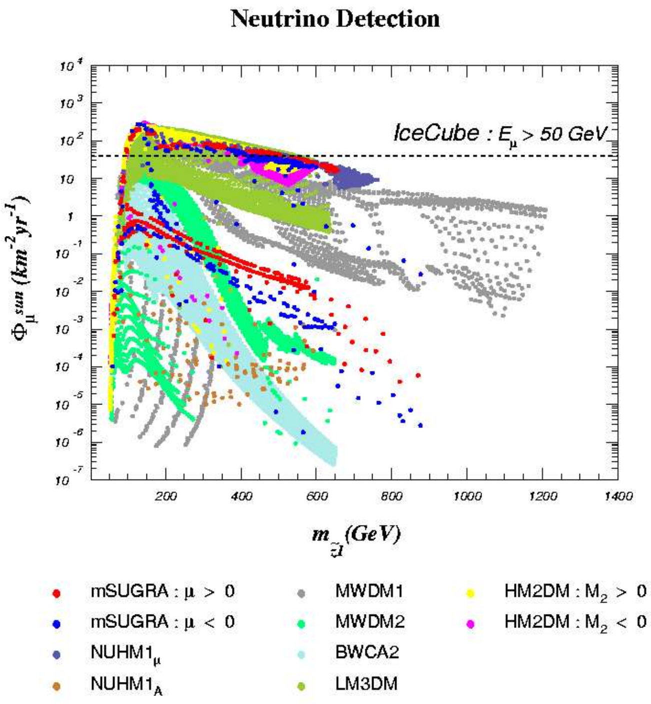
Implications for collider searches 2



- m_h *vs.* $m_{\tilde{t}_1}$
  - heavier $\tilde{t}_1$ squarks are correlated with larger values of m_h (due to top-Yukawa radiative corrections to m_h)
  - in many models with $m_A \gg M_Z$, then $h \simeq H_{\rm SM}$: the LEP2 lower bound of 114.1 GeV applicable
- $m_{\widetilde{W}_1}$ *vs.* $m_{\tilde{\tau}_1}$
  - dashed lines: reach of ILC500 ($\sqrt{s} = 500$ GeV)
  - dotted lines: reach of ILC1000 ($\sqrt{s} = 1000$ GeV)

Implications for direct/indirect(neutrino) DM detection

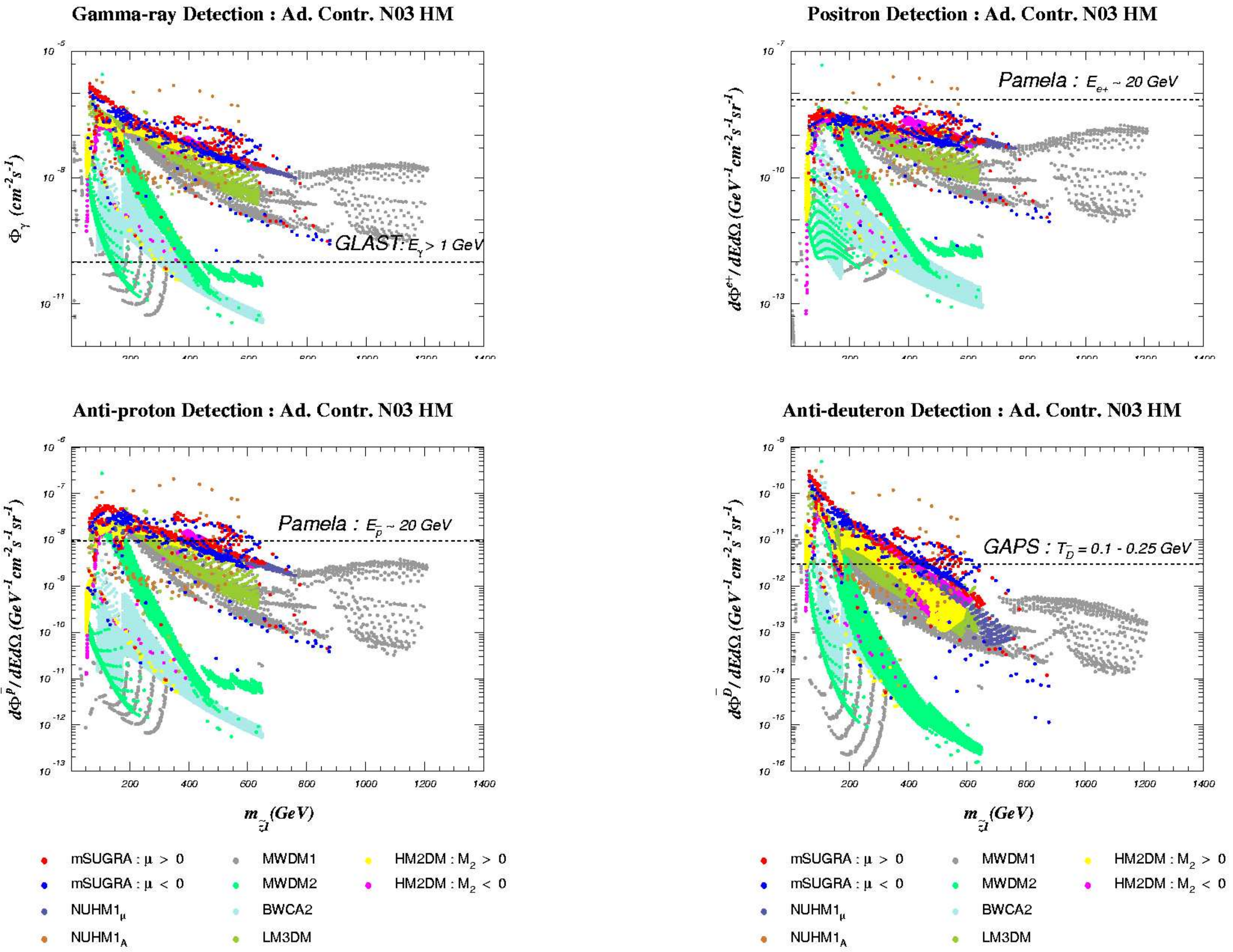




- models with WTN within reach of next generation of detectors
- models adjusted masses to get WMAP value below sensitivities of detectors

- muon fluxes from neutralino annihilation in the solar core to ν_μ states
- main contribution comes from Z-exchange ← enhanced if neutralino has high higgsino content

Implications for indirect(γ-ray, antiparticle) DM detection



Conclusions

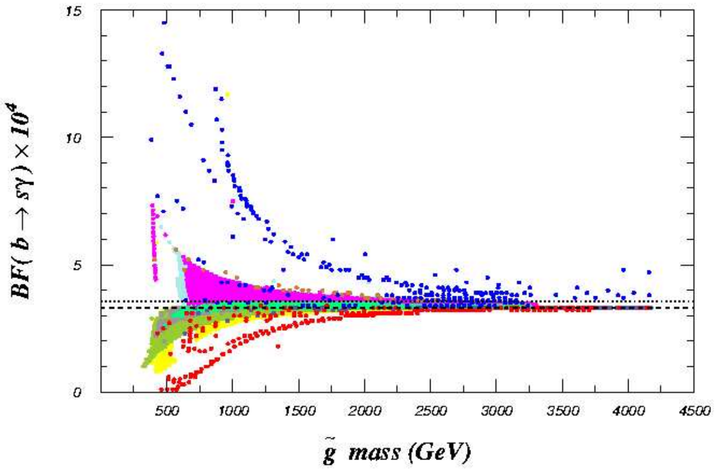
1. ⋆ WTN occurs *only* in FP/HB region in mSUGRA (MHDM: $m_{\tilde{q}} >> m_{\tilde{Z}_1, \tilde{W}_1, \tilde{g}}$).
   But, in relic-density-consistent models, easily get WTN with $m_{\tilde{q}} \sim m_{\tilde{g}}$
   ⋆ Higgs funnel enhancement is *only* for very large tanβ values in mSUGRA.
   But, in non-universal Higgs mass models, we have Higgs funnel for any tanβ value

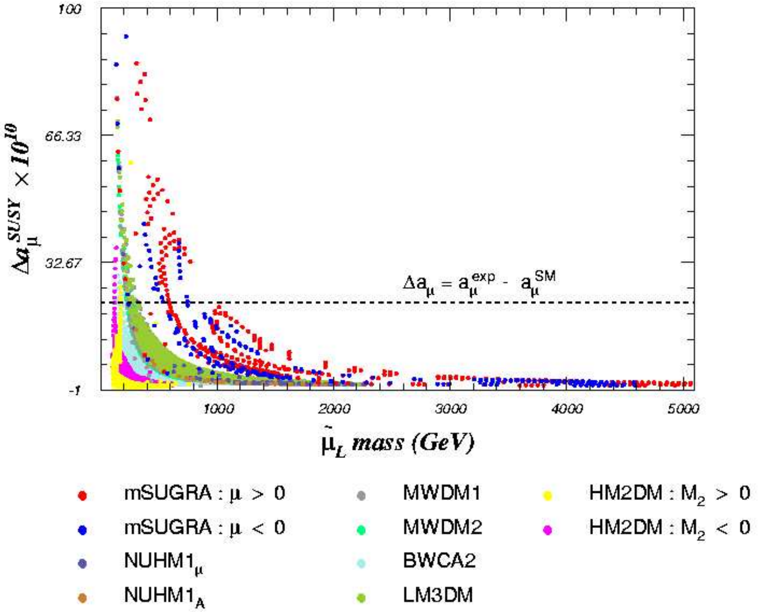
2. In many relic-density-consistent models, $\tilde{Z}_2 - \tilde{Z}_1$ mass gap $< M_Z$
   → 2-body decay modes kinematically closed
   → 3-body decay modes open ⇒ at least one dilepton mass edge detectable at LHC
   → location of dilepton mass edge is clean signature of SUSY models

3. ⋆ $m_{\tilde{q}} = m_{\tilde{g}}$, $m_{\tilde{q},\tilde{g}}$ <3100 GeV for most relic-density-consistent models
   → implies SUSY signals at LHC
   ⋆ $m_{\tilde{\tau}}$ <500 GeV for LM3DM
   → accessible at ILC with $\sqrt{s}$=1 TeV

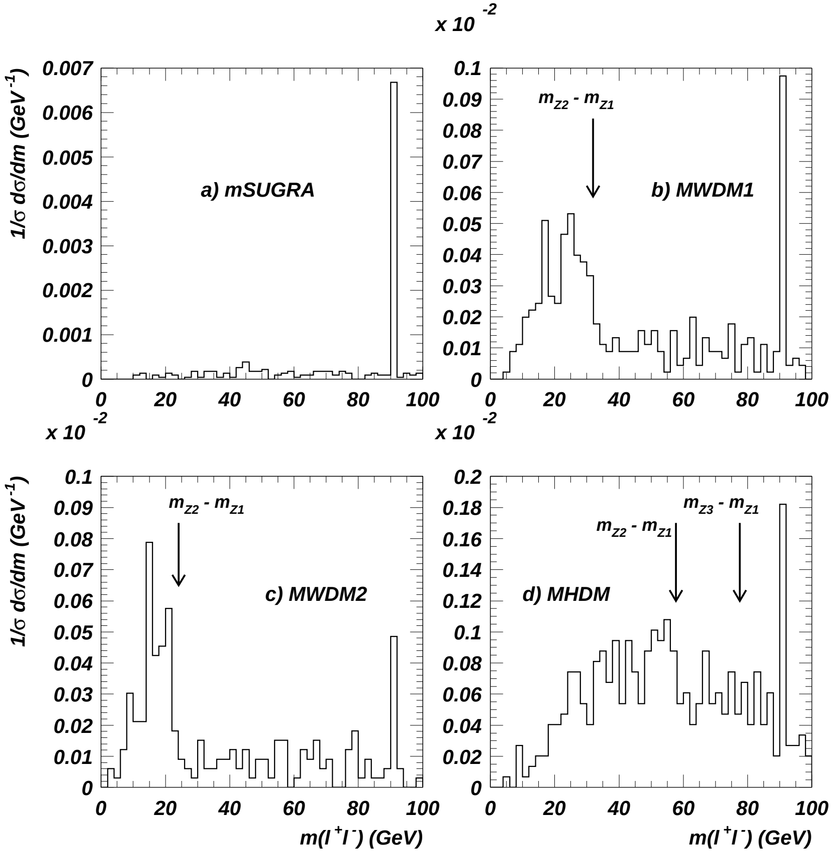
4. In WTN models,
   ⋆enhanced annihilation rates enhance direct DM detection rates
   ⋆in many cases, muon neutrino signals accessible at IceCube
   ⋆indirect DM searches in galactic halo into gamma rays and anti-matter elevated; large uncertainties associated with unknown galactic DM density profile

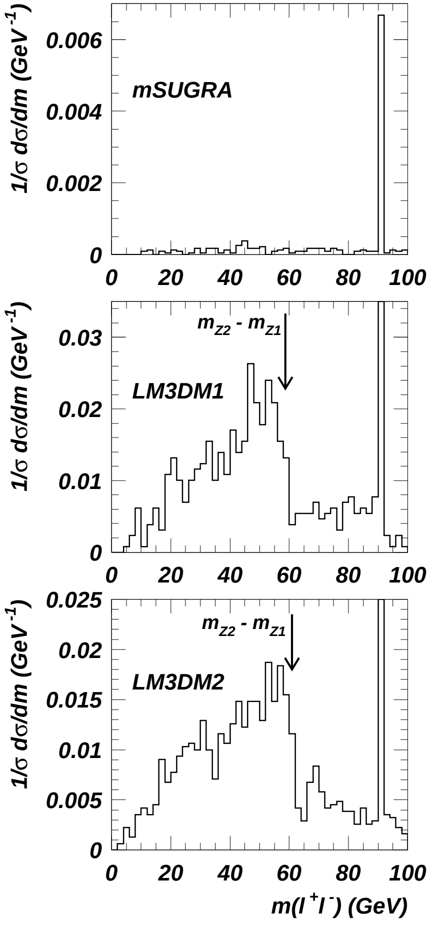
Implications for $BF(b \to s\gamma)$ and $(g-2)_\mu$





Dilepton Distribution at LHC





- mSUGRA : sharp peak at $m(l^+l^-) \sim M_Z$ from $\tilde{Z}_2 \longrightarrow \tilde{Z}_1 Z^0$ decays
- NUGM : Z^0 peak from $\tilde{Z}_3$, $\tilde{Z}_4$, $\tilde{W}_2$ decays + continuum distribution $m(l^+l^-) < m_{\tilde{Z}_2} - m_{\tilde{Z}_1}$

MSSM RGEs

$$\frac{dm^2_{H_u}}{dt} = \frac{2}{16\pi^2}\left(-\frac{3}{5}g_1^2M_1^2 - 3g_2^2M_2^2 + \frac{3}{10}g_1^2S + 3f_t^2X_t\right)$$

$$\frac{dm^2_{H_d}}{dt} = \frac{2}{16\pi^2}\left(-\frac{3}{5}g_1^2M_1^2 - 3g_2^2M_2^2 - \frac{3}{10}g_1^2S + 3f_b^2X_b + f_\tau^2X_\tau\right)$$

$$\frac{dm^2_{Q_3}}{dt} = \frac{2}{16\pi^2}\left(-\frac{1}{15}g_1^2M_1^2 - 3g_2^2M_2^2 - \frac{16}{3}g_3^2M_3^2 + \frac{1}{10}g_1^2S + f_t^2X_t + f_b^2X_b\right)$$

$$\frac{dm^2_{\tilde{t}_R}}{dt} = \frac{2}{16\pi^2}\left(-\frac{16}{15}g_1^2M_1^2 - \frac{16}{3}g_3^2M_3^2 - \frac{2}{5}g_1^2S + 2f_t^2X_t\right)$$

$$\frac{dm^2_{\tilde{b}_R}}{dt} = \frac{2}{16\pi^2}\left(-\frac{4}{15}g_1^2M_1^2 - \frac{16}{3}g_3^2M_3^2 + \frac{1}{5}g_1^2S + 2f_b^2X_b\right)$$

$$\frac{dm^2_{L_3}}{dt} = \frac{2}{16\pi^2}\left(-\frac{3}{5}g_1^2M_1^2 - 3g_2^2M_2^2 - \frac{3}{10}g_1^2S + f_\tau^2X_\tau\right)$$

$$\frac{dm^2_{\tilde{\tau}_R}}{dt} = \frac{2}{16\pi^2}\left(-\frac{12}{5}g_1^2M_1^2 + \frac{3}{5}g_1^2S + 2f_\tau^2X_\tau\right)$$

$$S = m^2_{H_u} - m^2_{H_d} + Tr\left[\mathbf{m}_Q^2 - \mathbf{m}_L^2 - 2\mathbf{m}_U^2 + \mathbf{m}_D^2 + \mathbf{m}_E^2\right]$$

where $t = \log(Q)$, $f_{t,b,\tau}$ are the t, b and τ Yukawa couplings, and

$$\begin{aligned} X_t &= m^2_{Q_3} + m^2_{\tilde{t}_R} + m^2_{H_u} + A^2_t \\ X_b &= m^2_{Q_3} + m^2_{\tilde{b}_R} + m^2_{H_d} + A^2_b \\ X_\tau &= m^2_{L_3} + m^2_{\tilde{\tau}_R} + m^2_{H_d} + A^2_\tau \end{aligned}$$

Some Benchmark Cases: non-universal scalar mass models

parameter	mSUGRA	NMH	NUHM1$_\mu$	NUHM1$_A$	HS
special	—	$m_0(1,2)$	m_ϕ	m_ϕ	δ_H
value	—	54	549	-728	-1.36
μ	385.1	386.5	105.8	748.5	269.3
$m_{\tilde{g}}$	729.7	722.1	731.4	733.4	728.9
$m_{\tilde{u}_L}$	720.8	658.4	724.3	720.5	720.1
$m_{\tilde{t}_1}$	523.4	526.5	484.1	624.5	505.8
$m_{\tilde{b}_1}$	656.8	659.8	642.2	689.5	645.4
$m_{\tilde{e}_L}$	364.5	216.2	364.8	365.8	373.4
$m_{\tilde{e}_R}$	322.3	128.9	322.5	321.9	301.8
$m_{\tilde{\tau}_1}$	317.1	317.6	317.8	316.4	299.3
$m_{\widetilde{W}_2}$	411.7	412.7	264.7	754.8	321.1
$m_{\widetilde{W}_1}$	220.7	219.5	91.1	234.9	196.6
$m_{\widetilde{Z}_2}$	220.6	219.4	117.4	234.5	198.1
$m_{\widetilde{Z}_1}$	119.2	118.4	69.0	121.5	115.4
m_A	520.3	521.9	584.5	268.5	279.0
m_{H+}	529.8	531.4	593.8	281.6	292.0
m_h	110.1	110.1	109.8	110.5	109.8
$\Omega_{\widetilde{Z}_1}h^2$	1.1	0.10	0.11	0.11	0.10
$\sigma_{SI}(\widetilde{Z}_1 p)$	2.1×10^{-9} pb	2.1×10^{-9} pb	7.8×10^{-8} pb	1.2×10^{-9} pb	2.7×10^{-8} pb
$R_{\tilde{H}}$	0.15	0.14	0.84	0.06	0.26

Some Benchmark Cases: non-universal gaugino mass models

parameter	mSUGRA	MWDM	BWCA	LM3DM	HM2DM
special	—	$M_1(M_{GUT})$	$M_1(M_{GUT})$	$M_3(M_{GUT})$	$M_2(M_{GUT})$
value	—	490	-480	160	900
μ	385.1	385.9	376.6	185.3	134.8
$m_{\tilde{g}}$	729.7	729.9	731.7	420.2	736.4
$m_{\tilde{u}_L}$	720.8	721.2	722.0	496.9	901.8
$m_{\tilde{u}_R}$	702.7	708.9	709.9	467.0	696.3
$m_{\tilde{t}_1}$	523.4	526.5	536.3	312.2	394.3
$m_{\tilde{b}_1}$	656.8	656.0	658.9	443.2	686.4
$m_{\tilde{e}_L}$	364.5	371.5	371.4	366.1	669.3
$m_{\tilde{e}_R}$	322.3	353.3	352.2	322.6	321.3
$m_{\widetilde{W}_2}$	411.7	412.4	404.5	282.9	719.7
$m_{\widetilde{W}_1}$	220.7	220.8	220.0	152.5	136.5
$m_{\widetilde{Z}_2}$	220.6	223.2	219.2	163.6	142.3
$m_{\widetilde{Z}_1}$	119.2	194.6	201.7	105.5	94.8
m_A	520.3	525.9	518.6	398.3	670.7
m_{H^+}	529.8	535.3	528.1	408.7	679.8
m_h	110.1	110.2	109.8	106.0	111.9
$\Omega_{\widetilde{Z}_1}h^2$	1.1	0.10	0.10	0.10	0.10
$\sigma_{SI}(\widetilde{Z}_1 p)$	2.1×10^{-9} pb	1.5×10^{-8} pb	3.1×10^{-11} pb	7.2×10^{-8} pb	3.4×10^{-8} pb
$R_{\tilde{H}}$	0.15	0.25	0.16	0.50	0.67